

GROUPS OF GALAXIES IN THE TWO MICRON ALL SKY REDSHIFT SURVEY

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ABSTRACT

We present the results of applying a percolation algorithm to the initial release of the Two Micron All Sky Survey Extended Source Catalog, using subsequently measured redshifts for almost all of the galaxies with $K < 11.25$ mag. This group catalog is based on the first near-IR all-sky flux-limited survey that is complete to $|b| = 5^\circ$. We explore the dependence of the clustering on the length and velocity scales involved. The paper describes a group catalog, complete to a limiting redshift of 10^4 km s^{-1} , created by maximizing the number of groups containing three or more members. A second catalog is also presented, created by requiring a minimum density contrast of $\delta\rho/\rho \geq 80$ to identify groups. We identify known nearby clusters in the catalogs and contrast the groups identified in the two catalogs. We examine and compare the properties of the determined groups and verify that the results are consistent with the UZC-SSRS2 and northern CfA redshift survey group catalogs. The all-sky nature of the catalog will allow the development of a flow-field model based on the density field inferred from the estimated cluster masses.

Subject headings: galaxies: clusters: general — galaxies: distances and redshifts — large-scale structure of universe

Online material: color figures, machine-readable tables

1. INTRODUCTION

The Two Micron All Sky Survey (2MASS) began in the early 1990s with the purpose of mapping the Milky Way and nearby universe. Previous all-sky surveys suffered from a variety of selection effects, many of them inconsistent across the sky. Observations at optical wavelengths suffer from severe extinction at low Galactic latitudes, motivating work using surveys conducted in the infrared. Many *IRAS*-selected galaxy samples have been investigated as tracers of the galaxy density field, e.g., Strauss et al. (1992) and Fisher et al. (1995) based on the 1.9 and 1.2 Jy samples, respectively, and Branchini et al. (1999) based on the PSCz catalog (Saunders et al. 2000); however, these samples are based on fluxes in the far-infrared and miss many early-type galaxies, and thus underestimate the total galaxy number density. Even though the *IRAS*-selected samples are not biased by extinction, they still suffer from confusion in high-density regions.

Since galaxies' spectra peak at $\sim 1.6 \mu\text{m}$, a survey in the near-infrared is optimized for their detection at the flux limit of the survey. By sampling uniformly over the entire celestial sphere in the J ($1.25 \mu\text{m}$), H ($1.65 \mu\text{m}$), and K ($2.16 \mu\text{m}$) bands, 2MASS has been designed to maximize the number of galaxies detected at a specified flux limit, producing the most complete all-sky survey performed to date (2MASS; Skrutskie et al. 2006).

Two catalogs, complete to $K \sim 13.5$ mag, were released in early 2003 (see the 2MASS explanatory supplement [Cutri et al. 2003] and Jarrett et al. 2000): a point-source catalog with 470,992,970

entries and an extended source catalog (XSC) with 1,647,559 objects classified as galaxies. Although designed for completeness down to low Galactic latitudes, the 2MASS XSC still suffers from confusion near the Galactic plane. The 2MASS Redshift Survey (2MRS; Huchra et al. 2005a, 2005b) uses the XSC as its input master list and aims to produce an all-sky, (extinction-corrected) flux-limited redshift catalog that will eventually be complete to $K = 13.0$ mag above $|b| = 5^\circ$. The 2MRS is currently 99.9% complete to $K = 11.25$ mag and $|b| > 5^\circ$.²

In this paper, we create a redshift-limited catalog of groups, uniformly sampled from the entire sky. By assuming the identified groups are virialized systems, we are able to provide estimates of the group masses, avoiding the necessity of assuming an intrinsic mass-to-light ratio. The local universe is sufficiently inhomogeneous at the scales in question that the dynamics due to our interactions with nearby groups are nonnegligible. Due to the nature of the all-sky group catalog presented here, we will now be able to estimate the local density field due to baryonic matter in the local universe. A flow-field model produced from this catalog can be used in conjunction with observations in order to answer the question of whether baryonic matter is a genuine tracer of dark matter.

The creation of group catalogs is not a new concept; however, the methods employed in developing these catalogs have evolved with the enhancements in instrumentation. Early group catalogs were based on limited or subjective data (e.g., de Vaucouleurs 1975), associating members based on similarity in apparent magnitude,

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² See 2MRS data release (J. P. Huchra et al. 2007, in preparation).

positional coincidence, and (if available) redshift. Turner & Gott (1976) proposed a method that identifies regions in which the surface number density on the sky is enhanced, creating group catalogs from two-dimensional data. This technique suffers because the typical angular separation of galaxies in a group will vary with distance, and thus nearby groups with large angular radii will not be identified. Furthermore, when applied to flux-limited surveys, this method will identify different groups for different limiting fluxes.

More recently, the use of objective algorithms to identify groups based on both their position on the sky and in redshift space has become widely accepted (e.g., Huchra & Geller 1982; Geller & Huchra 1983; Ramella et al. 1997, 2002; Diaferio et al. 1999; Giuricin et al. 2000), using methods designed to find the same groups regardless of the limiting magnitude of the sample. The applicability of a particular group-finding algorithm depends on the properties of the sample in question. For example, Marinoni et al. (2002) show that the Vornoi-Delauney method successfully reproduces the distribution of groups in velocity dispersion in a mock sample based on the Deep Extragalactic Evolutionary Probe (DEEP2) Redshift Survey (Davis et al. 2003); this method is adapted by Gerke et al. (2005) for application to the DEEP2 sample. The SDSS team developed an algorithm (C4; Miller et al. 2005) that searches for groups in three space dimensions as well as four photometric colors. Kochanek et al. (2003) used a matched filter algorithm to study clusters in 2MRS at the 89% completeness level. Yang et al. (2005) have developed a halo-based group finder and successfully applied it to the 2dFGRS sample (Merchán & Zandivarez 2002; Eke et al. 2004). The same technique has been applied to the SDSS by Weinmann et al. (2006).

In this paper we apply a variable linking-length percolation (also commonly referred to as a *friends-of-friends*) algorithm (Huchra & Geller 1982, hereafter HG82) to determine the groups present in 2MRS. The velocity dispersion within the groups will allow estimates of the virial masses of the groups, thus providing a method to trace the density field associated with luminous matter in the local universe.

We begin with an outline of the group-identification algorithm in § 2. We discuss the modifications made to the data sample prior to the application of the algorithm in § 3. Section 4 presents a discussion on the choice of parameters used in the group-identification algorithm. The group catalogs and their properties are discussed in § 5, and we summarize our conclusions and discuss the potential applications of the catalogs in § 6.

2. GROUP-IDENTIFICATION ALGORITHM

We use the algorithm described in HG82 to identify groups of galaxies in the $K < 11.25$ mag version of the 2MRS catalog. The procedure is outlined briefly below.

We compare each galaxy in the catalog with its neighboring galaxies; for each pair of galaxies a linking length $D_L(V_{\text{avg}})$ is computed that depends on the average redshift of the galaxies, V_{avg} . Given two galaxies with an angular separation θ , we ask whether their projected separation, $D_{12} = \sin(\theta/2)V_{\text{avg}}/H_0$, is less than $D_L(V_{\text{avg}})$. If this is true, and the difference in redshift, $V_{12} = |V_1 - V_2|$, is less than some linking velocity V_L , then we identify both galaxies with the same group. D_L is defined as follows:

$$D_L = D_0 \left[\frac{\int_{-\infty}^{M_{12}(V_{\text{avg}})} \Phi(M) dM}{\int_{-\infty}^{M_{\text{lim}}} \Phi(M) dM} \right]^{-1/3}, \quad (1)$$

where

$$M_{12}(V_{\text{avg}}) = m_{\text{lim}} - 25 - 5 \log(V_{\text{avg}}/H_0).$$

Here $\Phi(M)$ represents the differential galaxy luminosity function for the sample and D_0 the projected separation (in Mpc) at some chosen fiducial redshift V_F . $M_{\text{lim}} = M_{12}(V_F)$ is a constant for a given V_F , and m_{lim} is the apparent-magnitude limit of the sample.

This scaling of the linking length compensates for the bias that would otherwise be introduced due to the variation in sampling of the luminosity function with redshift. There is much debate on how and whether or not to scale V_L (e.g., HG82; Nolthenius & White 1987; Frederic 1995a, 1995b). If one assumes uniform density, simple scaling arguments show that the velocity is simply proportional to the radius, suggesting that V_L should be scaled in the same manner as D_L . Such a scaling would include unwanted interlopers at large values of V_L and thus introduce an unwanted correlation between velocity dispersion and redshift. The density profiles of galaxy clusters, however, are usually better described by the isothermal-sphere approximation; in this case the velocity dispersion is independent of the size of the cluster. It follows therefore that by setting V_L to a reasonable fixed value, we will minimize the number of interlopers, but not bias the algorithm against finding distant groups. Hereafter we set

$$V_L = V_0. \quad (2)$$

The choice of D_0 determines the minimum density contrast of identified groups, which can be estimated using (see HG82)

$$\frac{\delta\rho}{\rho} = \frac{3}{4\pi D_0^3} \left[\int_{-\infty}^{M_{\text{lim}}} \Phi(M) dM \right]^{-1} - 1. \quad (3)$$

3. THE SAMPLE

The first available sample of the 2MRS galaxy catalog contains positions, redshifts, and magnitudes for 23,090 galaxies selected from the XSC. The targets were by selected by introducing a cut on the corrected magnitudes of objects in the XSC of $K < 11.25$ mag (the apparent magnitudes had previously been corrected for extinction using the dust maps of Schlegel et al. 1998). This catalog is complete, bar 40 galaxies, for Galactic latitudes $|b| > 10^\circ$ between Galactic longitudes 330° and 30° , and $|b| > 5^\circ$ for other longitudes.

Below we discuss a simple flow-field model applied to provide improved estimates of the distances to the galaxies (see § 3.1). In § 3.2 we discuss the method used to populate the Galactic plane with random galaxies to prevent any artifacts arising from the significantly reduced observed number density of galaxies behind the plane. We briefly discuss the assumed form of the luminosity function of the sample in § 3.3, and in § 3.4 we consider the completeness of the sample in redshift space.

3.1. Distance Estimates

Locally, redshifts do not provide a reliable indication of distance because of distortions to the local velocity field due to infall onto concentrations of mass. Although the clustering algorithm is independent of the observer's frame of reference, it is essential to have reasonable estimates of the distances to the galaxies in order to compute the linking parameters D_{12} and D_L , as well to accurately estimate the luminosities of the galaxies.

We apply the basic flow-field model described in Mould et al. (2000) to account for the local distortions to the velocity field.

This prescription first corrects the reference frame to the Local Group frame (Yahil et al. 1977, corroborated by the more recent work of Courteau & van den Bergh 1999), then adjusts the redshift-inferred distances of galaxies near Virgo, Shapley, and the Great Attractor (GA) region as follows: All galaxies within 12° of the center of Virgo with heliocentric redshifts less than 2500 km s^{-1} are placed at the redshift of Virgo (plus a random velocity, drawn from a Gaussian distribution with a standard deviation of 20 km s^{-1} , to avoid artifacts in the group properties occurring from galaxies with identical redshifts). All galaxies within 10° and 2000 km s^{-1} of the GA are placed at the redshift of the GA (plus scatter), and all galaxies within 12° and 3000 km s^{-1} of Shapley are placed at the redshift of Shapley (plus scatter). The corrected velocities are then used in place of the heliocentric velocities when computing distances only. To infer the distances we assume Hubble's law is valid to the completeness limit of the 2MRS catalog,³ using a Hubble constant $H_0 = 100 h \text{ km s}^{-1} \text{ Mpc}^{-1}$, where we assume $h = 0.73$ when a specific value is required. This value is chosen based on the 3 year *WMAP* results (Spergel et al. 2006), $h = 0.73 \pm 0.03$. In the very local universe (i.e., where corrected distances are less than $3 h^{-1} \text{ Mpc}$), we give galaxies an indicative distance of $3 h^{-1} \text{ Mpc}$. The velocities used in computation of V_{12} , etc., are the heliocentric velocities reported in the 2MRS catalog.

3.2. Filling in the Galactic Plane

The 2MRS catalog is currently incomplete near the Galactic plane ($|b| < 10^\circ$ between Galactic longitudes 330° and 30° , and $|b| < 5^\circ$ for other longitudes). With a significantly reduced number density of galaxies observed behind the Galactic plane, a structure that spans the plane will not be identified by the clustering algorithm. Similarly, structures that are visible in part above or below the plane may not be identified as groups, and even if they are, a bias will be introduced in the number density of groups with centers just above or below the plane. Any flow-field model derived from such a group catalog will suffer from these biasing effects; we therefore attempt to minimize these effects by randomly populating the sample to enhance the galaxy number density behind the Galactic plane to reflect that observed above and below it.

We follow a method similar to that of Yahil et al. (1991); this method produces similar results to the more involved Wiener reconstruction⁴ discussed in Lahav et al. (1994). We first divide the catalog into bins spanning 10° in Galactic longitude and $10 h^{-1} \text{ Mpc}$ in distance. For Galactic longitudes ranging from 330° to 30° (masking the bulge), we now consider bins further bounded by the lines $|b| = 10^\circ$. Sampling from the adjacent bins ($10^\circ < |b| < 20^\circ$), we populate the bulge with N galaxies drawn at random from the galaxies in adjacent bins; these galaxies are placed at random latitudes, and a normal scatter of 20 km s^{-1} is introduced in the velocity to prevent artifacts in the group properties arising due to galaxies at identical redshifts. N is calculated by drawing a random normal deviate from a distribution with a mean equal to the number of galaxies in the two adjacent bins (above and below the plane), then subtracting the number of galaxies already present within the bin. For other Galactic longitudes, the latitudes $5^\circ < |b| < 15^\circ$ are used to populate the bins with

$|b| < 5^\circ$.⁵ The catalog, before and after population of the Galactic plane, is shown in Figure 1. The population of the plane generated an additional 2076 galaxies.

3.3. Luminosity Function

The *K*-band luminosity function utilized in the algorithm is parameterized in terms of a function of the form (Schechter 1976)

$$\Phi(M) = 0.4 \ln(10) \Phi^* 10^{0.4(\alpha+1)(M^*-M)} \exp(-10^{0.4(M^*-M)}).$$

We use the values reported in Huchra et al. (2005a),

$$\alpha = -1.02, \quad M^* = -24.2, \quad \Phi^* = 1.08 \times 10^{-2} h^3 \text{ Mpc}^{-3}, \quad (4)$$

which have been computed using the galaxies in the 2MRS catalog with Galactic latitudes, $|b| > 10^\circ$.

3.4. Completeness in Redshift Space

Due to the nature of flux-limited surveys, the number density of galaxies observed at sufficiently high redshifts will tend toward zero. At these highest redshifts, the linking lengths used in the algorithm become so large that the majority of identified groups will likely be spurious. For the purposes of building a flow-field model, the groups at the highest redshifts will have the smallest affect on local dynamics; thus we choose to limit the group catalog to a redshift inside which the catalog is reasonably complete. Figure 2 shows the cumulative number of galaxies, $N(<D)$, as a function of (estimated) distance, D . The points have been fitted with a curve of the form

$$N(<D) = N_0 \left(\frac{\beta D}{[(\beta D)^b + S^b]^{1/b}} \right)^a, \quad (5)$$

where the best-fit parameters are $N_0 = 2.57 \times 10^4$, $a = 2.10$, $S = 106 \text{ Mpc}$, $\beta = 0.881$, and $b = 3.94$. The derivative of equation (5) represents the selection function of the survey, $N(D)$, where the number of galaxies observed with estimated distances between D and $D + dD$ is given by $N(D)dD$. We choose to cut the group catalog at the distance where the selection function falls to half its maximum, $D_{\text{cut}} = 140 \text{ Mpc}$, illustrated in Figure 2. The entire 2MRS $K < 11.25$ data set plus the galaxies generated in the population of the Galactic plane will be used to create the group catalog, but the catalog will then be truncated, excluding groups with mean estimated distances greater than D_{cut} .

4. PARAMETER CHOICES

In this section we justify the choice of linking parameters used in the construction of our group catalog. Any group catalog produced from the remaining data sample will contain minimal biasing effects at the highest and lowest redshifts, as well as minimal edge effects across the Galactic plane. There remains, however, a choice of the parameters that specify the minimum density contrast of detected groups. These parameters, D_0 and V_0 from equations (1) and (2) above, must be chosen in a somewhat arbitrary fashion.

There is no perfect choice of these parameters that will allow us to identify only groups that are gravitationally bound. In any

³ At the limiting redshifts of the galaxies analyzed, the difference between distances computed using a Λ CDM cosmology and simply assuming Hubble's law is less than 5%.

⁴ See Erdoğdu et al. (2006) and A. Rassat et al. (2007, in preparation) for discussions on the Wiener reconstruction of the 2MRS sample.

⁵ In this case, we set the mean of the normal distribution from which N is drawn to half the number of galaxies in the adjacent bins.

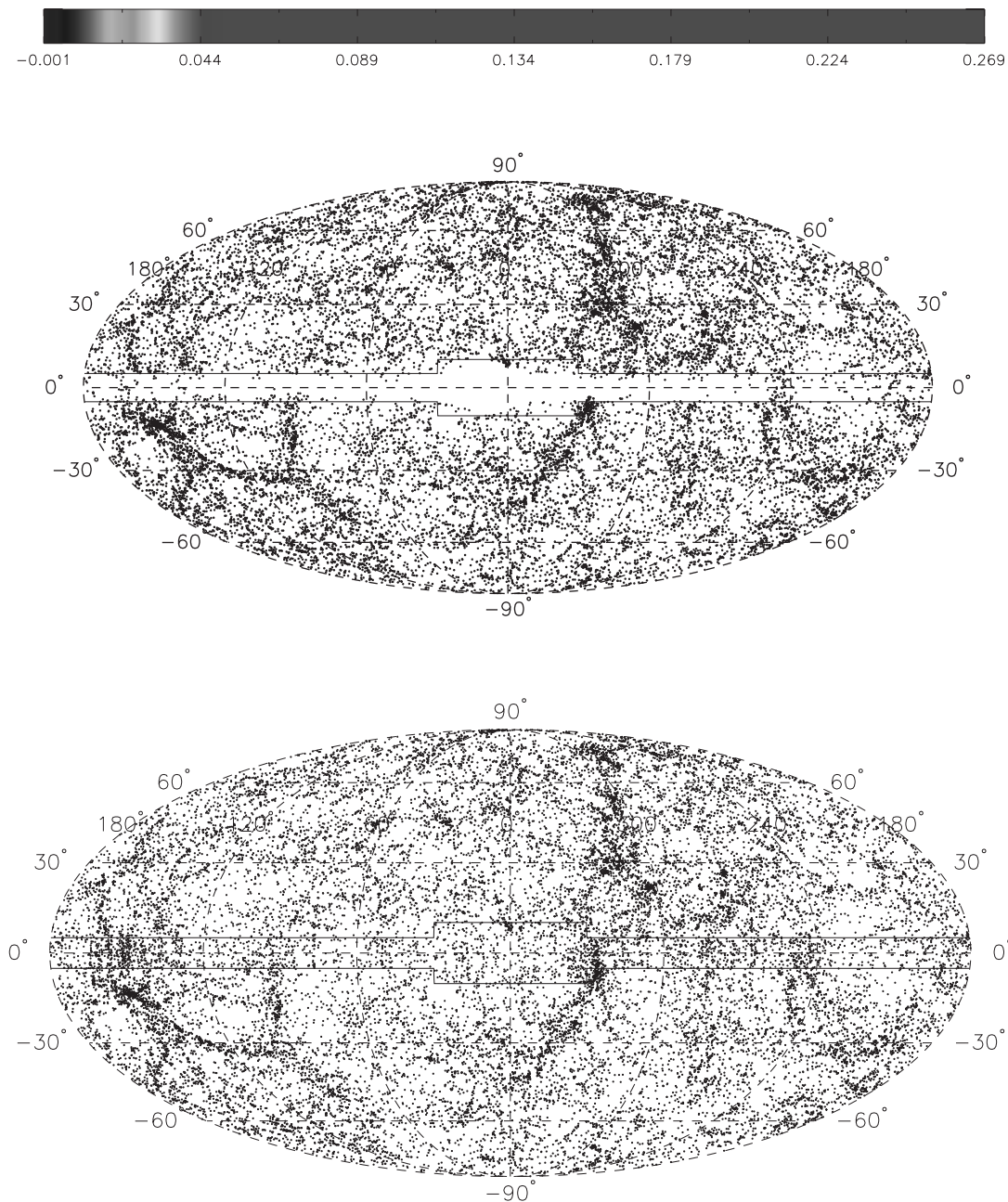


FIG. 1.—Galaxies in 2MRS catalog shown in a Mollweide projection in Galactic coordinates. *Top*: Catalog before the plane was populated. *Bottom*: Catalog including the addition of the randomly generated galaxies. The solid line indicates the region that was populated. (The color in the electronic edition indicates the measured redshift of the galaxy.) [See the electronic edition of the *Journal* for a color version of this figure.]

choice we make, some bound systems may be divided and unbound galaxies will be present in some groups. For very large values of both D_0 and V_0 , the algorithm will associate all of the galaxies into a single group. Conversely, if we choose sufficiently small D_0 or V_0 , the algorithm will divide substructures within real clusters into multiple systems (Ramella et al. 1997), eventually separating each galaxy into its own group. It is clear, therefore, that a suitable parameter choice will lie between these extreme cases. The method of choosing the specific values of the parameters must still remain arbitrary; in order to be able to infer properties of the universe (e.g., the matter density parameter) from the catalog, it is unwise to calibrate the algorithm using simulations based on a set of defined initial assumptions as this would bias

our results toward recovering these initial values.⁶ It is obvious that there will be a choice of parameters that lie between these extreme values that maximizes the *number* of groups produced. It is therefore reasonable to use a method of maximization to determine the choice of linking parameters, with no alternative method available that does not have similar shortcomings.

At this point, we must consider the size of the group we choose to maximize. We choose to ignore binaries in our definition of groups, as previous work has shown such systems identified using percolation algorithms to be unbound in the majority of cases (e.g., Diaferio et al. 1999). We consider the parameters obtained

⁶ See A. C. Crook et al. (2007a, in preparation) for further investigation.

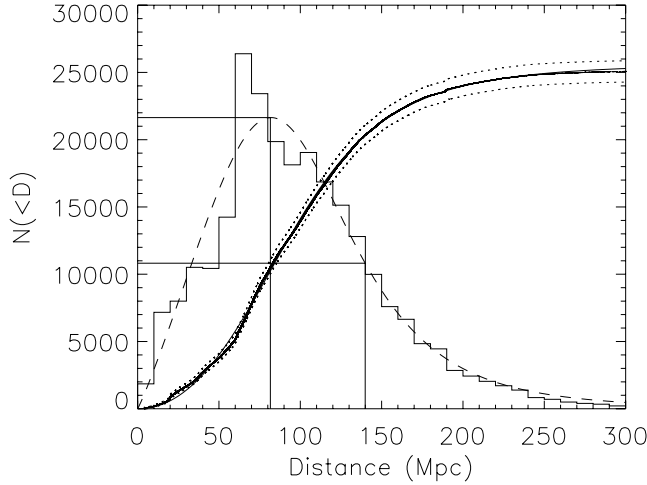


FIG. 2.—Selection function of the 2MRS survey. The data points (*small dots*) represent the cumulative number of galaxies as a function of (estimated) distance. The data are fit with a curve of the form of eq. (5) using a least-squares fit (*solid line*). The dotted lines show the 5σ errors from Poisson statistics. The data depart from the fit at the both smallest and largest distances shown on the plot. The dashed line shows the derivative of this curve (the selection function) in arbitrary units. The histogram contains the binned data shown in the same arbitrary units as the selection function for purposes of comparison only. The maximum and half-maximum values of the selection function are indicated.

when maximizing the number of groups of G or more members for $3 \leq G \leq 20$, as described below. We choose to set $V_F = 1000 \text{ km s}^{-1}$ following HG82. Figure 3 shows the number of groups containing three or more members in D_0 - V_0 space. In Figure 3a we explore the parameter space on the intervals $D_0 = [0, 10] \text{ Mpc}$, $V_0 = [0, 2000] \text{ km s}^{-1}$. We then attempt to maximize the number of groups obtained by the following method: we divide the region spanning $0 \rightarrow 10 \text{ Mpc}$ in D_0 and $0 \rightarrow 2000 \text{ km s}^{-1}$ in V_0 into a 9×9 grid and search for the combination of parameters that produces the largest number of groups. We then change the range of the D_0 and V_0 parameters spanned to coincide with a 3×3 grid (as far as possible) centered on the values of D_0 and V_0 that produced the largest number of groups. We divide this region into a 9×9 grid and iteratively repeat the procedure until the desired accuracy of the parameters is reached. This procedure is illustrated in Figure 3b.

We repeat this maximization procedure for all values of G between 3 and 20. The obtained value of D_0 rises gently with G , while the scatter in obtained values of V_0 increases rapidly with G . As observed in Figure 3a, the number of groups depends highly on D_0 , but is comparatively insensitive to V_0 . Since the velocity dispersion of a cluster is not expected to depend on the size of the cluster (see the discussion in § 2 above), it is not sensible to consider large values of V_0 (i.e., $> 1000 \text{ km s}^{-1}$) as this will introduce members that are not gravitationally bound and thus propagate errors into the mass estimates. The fraction of interlopers present in a group will also increase with both D_0 and V_0 ; thus the most sensible choice of parameters to minimize interlopers and reduce the scatter in V_0 corresponds to $G = 3$.

The maximum number of groups of three or more members (1538) is obtained for the values $(D_0, V_0) = (1.63 \pm 0.03 \text{ Mpc}, 399 \pm 8 \text{ km s}^{-1})$, corresponding to the density contrast $\delta\rho/\rho = 12$.

In an analysis of the northern CfA redshift survey⁷ (hereafter CfAN), Ramella et al. (1997) show that the group properties are

statistically stable for values of density contrasts $\delta\rho/\rho \geq 80$, where they scale V_L in a similar manner to D_L and choose $V_0 = 350 \text{ km s}^{-1}$. Diaferio et al. (1999) apply a similar choice of parameters to mock CfA surveys based on N -body simulations and conclude that 80% of groups with four or more members are true virialized systems, whereas 40% of triplets are not, confirming the hypothesis of Ramella et al. (1989). As Figure 3a shows minimal variation in the number of groups produced with changing V_0 compared to changing D_0 , the findings of Ramella et al. (1997) are applicable to this study. We will proceed to analyze the groups produced at both the values of D_0 and V_0 that maximize the number of groups of three or more members, as well as the values suggested by Ramella et al. (1997) (i.e., $D_0 = 0.89 \text{ Mpc}$, which corresponds to $\delta\rho/\rho = 80$,⁸ and $V_0 = 350 \text{ km s}^{-1}$).

5. GROUPS

We present the results of applying the group-finding algorithm (§ 2) to the 2MRS catalog subset (§ 3) using both pairs of parameters discussed in § 4 above. The group catalogs are presented in Tables 6–11 in the Appendix (see the electronic edition for the complete catalogs). We provide an overview of the catalogs in § 5.1 below, then discuss the identified groups and contrast the two catalogs in § 5.2. We present the properties of the obtained groups (§ 5.3) and discuss the reliability of the clustering algorithm (§ 5.4).

5.1. Overview

The catalog produced using the parameters $(D_0, V_0) = (1.63 \text{ Mpc}, 399 \text{ km s}^{-1})$ is presented in Table 6. These parameters produced the maximum number of groups of three or more galaxies and correspond to a density contrast $\delta\rho/\rho = 12$; this catalog will hereafter be referred to as the low-density-contrast (LDC) catalog. The catalog produced using the parameters $(D_0, V_0) = (0.89 \text{ Mpc}, 350 \text{ km s}^{-1})$ is presented in Table 7. These parameters correspond to the density contrast $\delta\rho/\rho = 80$; this catalog will hereafter be referred to as the high-density-contrast (HDC) catalog.

Figure 4 shows the positions and sizes of all groups in the two catalogs. Figures 4a and 4b show the groups in equatorial coordinates; Figures 4c and 4d show the groups in Galactic coordinates.

The maps are shown in Mollweide projection, which preserves the area of the structures on the surface of a sphere but distorts their shape, especially close to the poles. The plots show ellipses that have the properties of the groups discussed in § 5.3.2 below, and are only representative of the shape and size of the group. The Local Supercluster (LSC) has been clearly identified in Figure 4a as the large structure in the center of the figure. When applying the algorithm with the higher minimum density contrast, this structure is split into several constituents as shown in Figure 4b. (The same result is evident in Figs. 4c and 4d; however, the LSC encompasses the pole of the coordinate system in this case. The distortions due to the map projection are therefore enhanced in these plots, and the shape is less representative of the true shape of the LSC. The area occupied by the LSC in Fig. 4c is the same as that in Fig. 4d. The constituents that have been merged to form the LSC are clearly visible in Fig. 4d.)

There is an apparent enhancement in the number of groups with large angular sizes near the Galactic plane (see Figs. 4b and 4d). There are five groups shown as ellipses with major axes greater than 5° and with centers inside $|b| < 10^\circ$. Of these, two

⁷ This is a subset of the extended CfA redshift survey (de Lapparent et al. 1991; Geller & Huchra 1989; Huchra et al. 1990, 1995).

⁸ The smallest allowed density contrast is chosen to minimize the probability of splitting the richest systems.

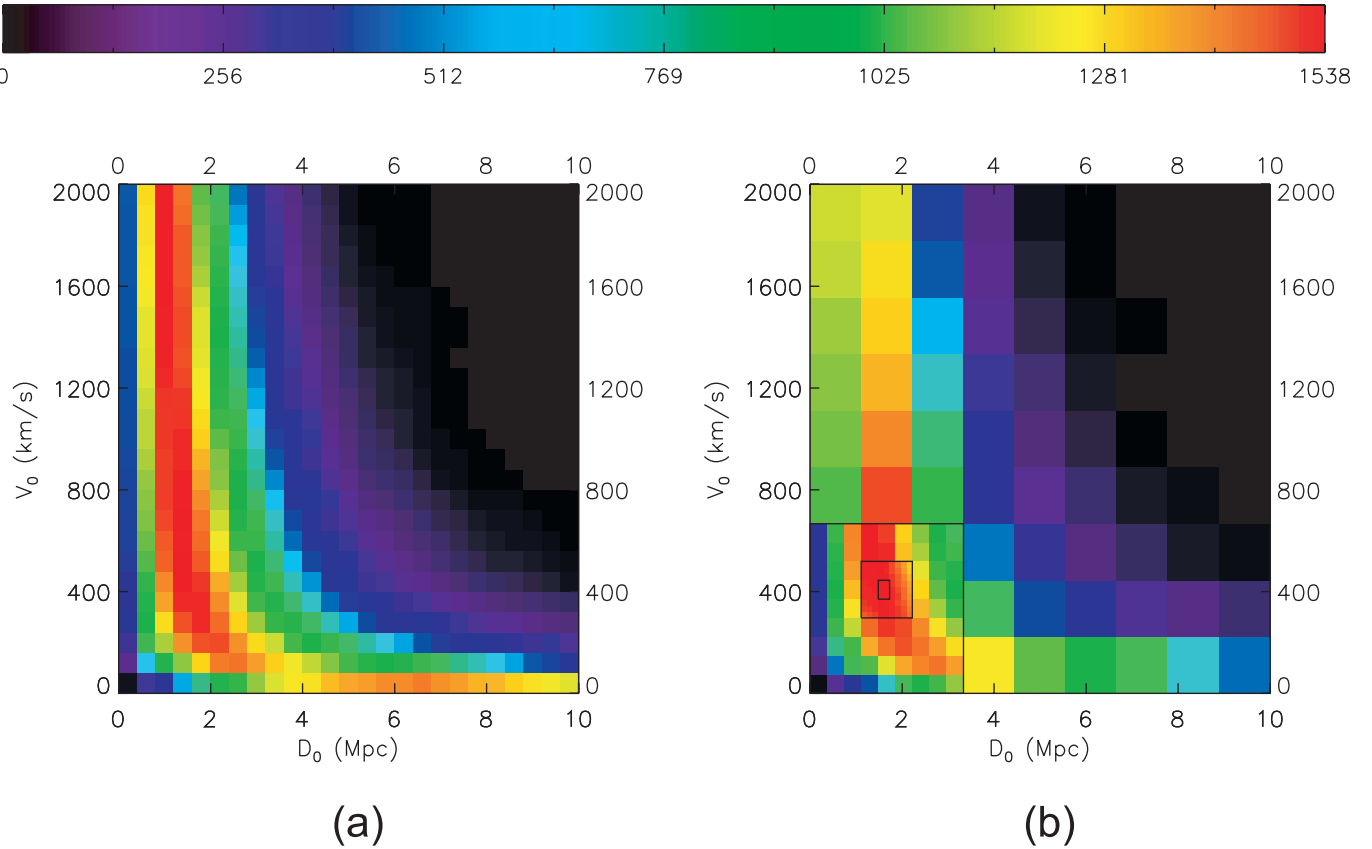


FIG. 3.—Number of groups of three or more galaxies obtained as a function of the parameters D_0 and V_0 . In (a), the clustering algorithm has been executed for each pair of parameters on a 25×25 grid; (b) contains a graphical representation of the execution of the maximization routine discussed in the § 4.

have only three members (of which two are genuine galaxies from 2MASS XSC), and only one out of the remaining three groups has more than 10% of its members randomly generated in the population of the plane; the remaining two are the only groups with more than five members; thus the large apparent sizes of all five groups in the figure are due to their proximity. We conclude that the observed enhancement is therefore not an artifact of the population of the plane.

Figure 5 shows the same groups as Figure 4; however, the area of each ellipse is proportional to the number of members in the group, rather than the square of the group's angular size.

The areas are normalized such that the major axis (see § 5.3.2) of the largest group appears as 75% of its true angular size. These figures are no longer dominated by the foreground groups that have the largest angular sizes, but show how the groups are distributed across the sky out to the redshift limit of the sample.

5.2. Identification and Overlap

Due to the nature of the group-finding algorithm, we expect to find that all of the galaxies assigned to groups in the HDC catalog will also be assigned to groups in the LDC catalog; however, the converse is not necessarily true. We consider all the galaxies assigned to groups in the HDC catalog and determine the correspondence between groups in the two catalogs. The six largest groups in the LDC catalog are plotted in Figure 6; the corresponding groups in the HDC catalog are also shown. The largest group in the LDC catalog contains 810 members (of which two were randomly generated in the population of the Galactic plane). This group is the result of merging several smaller groups, including the Virgo, NGC 3607, NGC 4105, IC 764, NGC 5746, NGC

3190, NGC 5846, and NGC 4038 clusters, as well as the M81 group, and corresponds to the LSC. In the HDC catalog, most of these groups have been identified individually; in fact Virgo has been split into two groups, containing 298 galaxies and 123 galaxies, respectively. Eridanus, Fornax I, Dorado, and the NGC 2280, NGC 1433, and NGC 2559 clusters merge to form the second largest group in the LDC catalog (302 members, including 22 simulated); again these were identified individually in the HDC catalog. A426 and A347 of the Perseus-Pisces supercluster make up the third largest group in the LDC catalog, containing 301 galaxies, but these were identified as two separate groups in the HDC catalog. Hydra (A1060) is identified as the fourth largest with 241 members, and Norma (A3627, the GA) is identified as the fifth with 217 members (42 of which were randomly generated). Centaurus (A3526) was identified as the sixth largest group (202 members). This exercise demonstrates that the correct choice of parameters used in the group algorithm is entirely dependent on the size of the structures that are desired.

The remaining identifications and correspondence between the catalogs are shown in the Appendix (Tables 8 and 9), where we consider only groups containing 25 or more members in the HDC catalog. Tables 10 and 11 contrast the group assignments of individual galaxies between the two catalogs. It may be surprising that the parameters chosen to maximize the total number of groups actually merge several of the large groups, hence apparently reducing the total number of groups. Although not obvious, this result is not unexpected because the larger linking length will allow many smaller groups to be identified that do not exist in the HDC catalog. The latter association is generating more groups than are removed by the merging of largest groups. It is likely that the groups

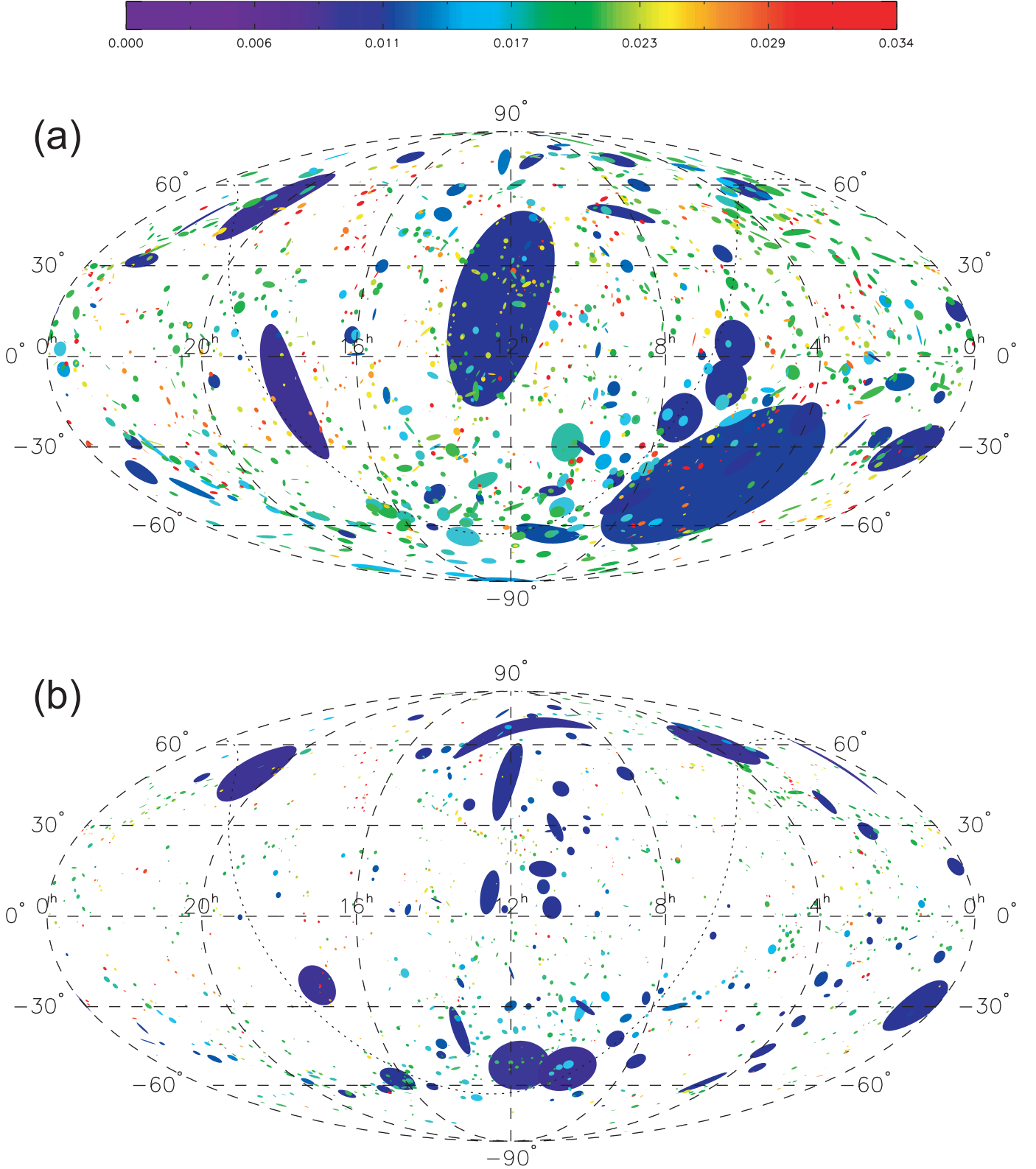


FIG. 4.—Groups (scaled by angular size) identified by clustering algorithm in 2MRS catalog. Panels *a* and *b* are shown using a Mollweide projection in equatorial coordinates, centered at $12^{\text{h}}00^{\text{m}}$. Panels *c* and *d* are shown using a Mollweide projection in Galactic coordinates. Each group is plotted as an ellipse with the major axis, position angle, and axis ratio of the group. The ellipses are transformed from the x - y coordinate system discussed in § 5.3.2 to the appropriate map projection. The color of the ellipse represents the group's mean redshift. The Galactic plane is shown by the dotted line in panels *a* and *b*. Panels *a* and *c* show the groups in the LDC catalog ($\delta\rho/\rho = 12$); the LSC has been clearly identified as the central structure. Panels *b* and *d* show the groups in the HDC catalog ($\delta\rho/\rho = 80$); we observe the effect of increasing the minimum density contrast required to identify groups: the LSC has been broken into several constituents. The distortions created by the map projection are maximized near the poles and enhanced for larger structures. The seemingly strange shape of the LSC in panel *c* is the result of mapping an ellipse in the x - y coordinate system onto this projection, and is only partly representative of the true shape of the structure.

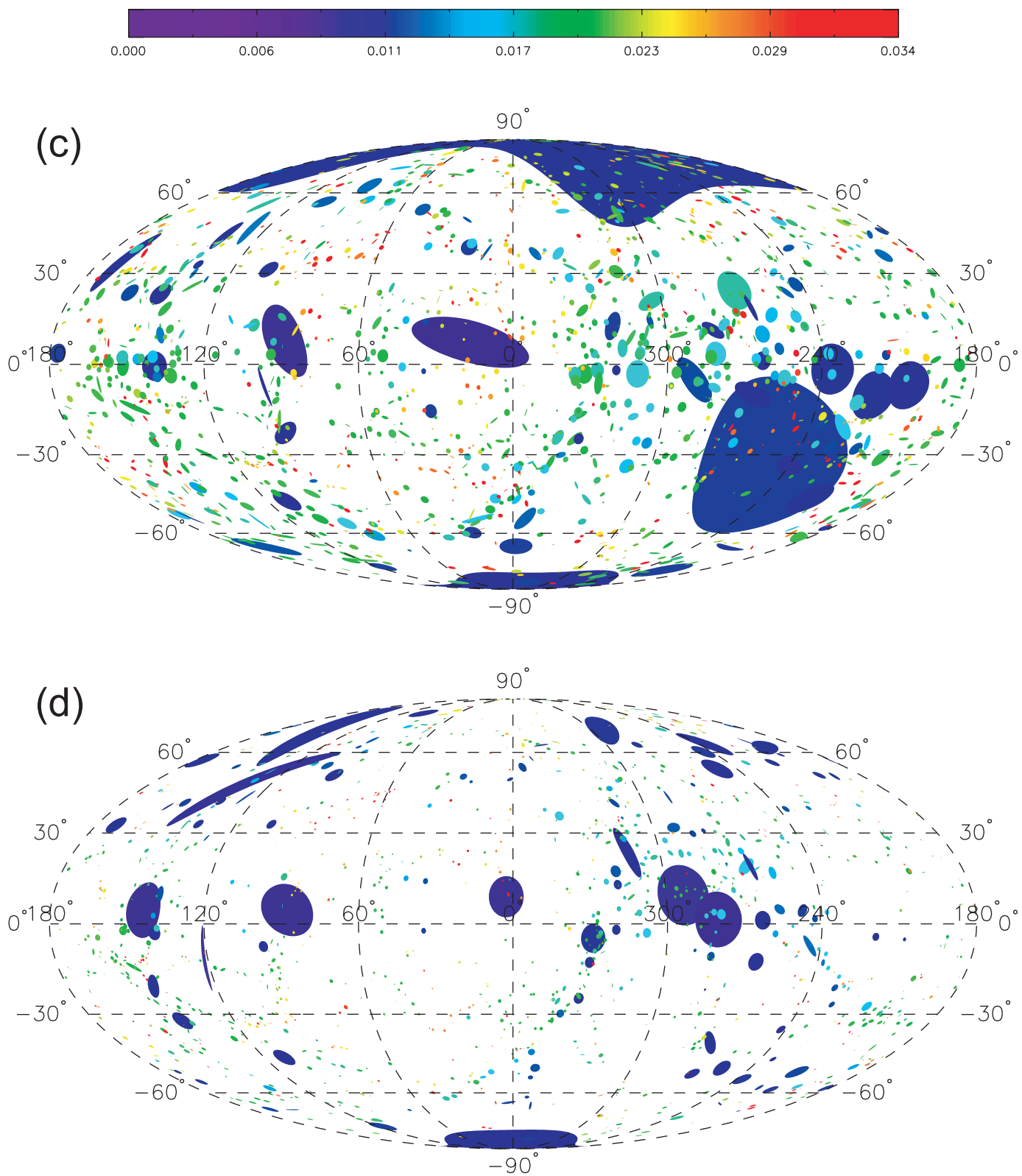


FIG. 4.— *Continued*

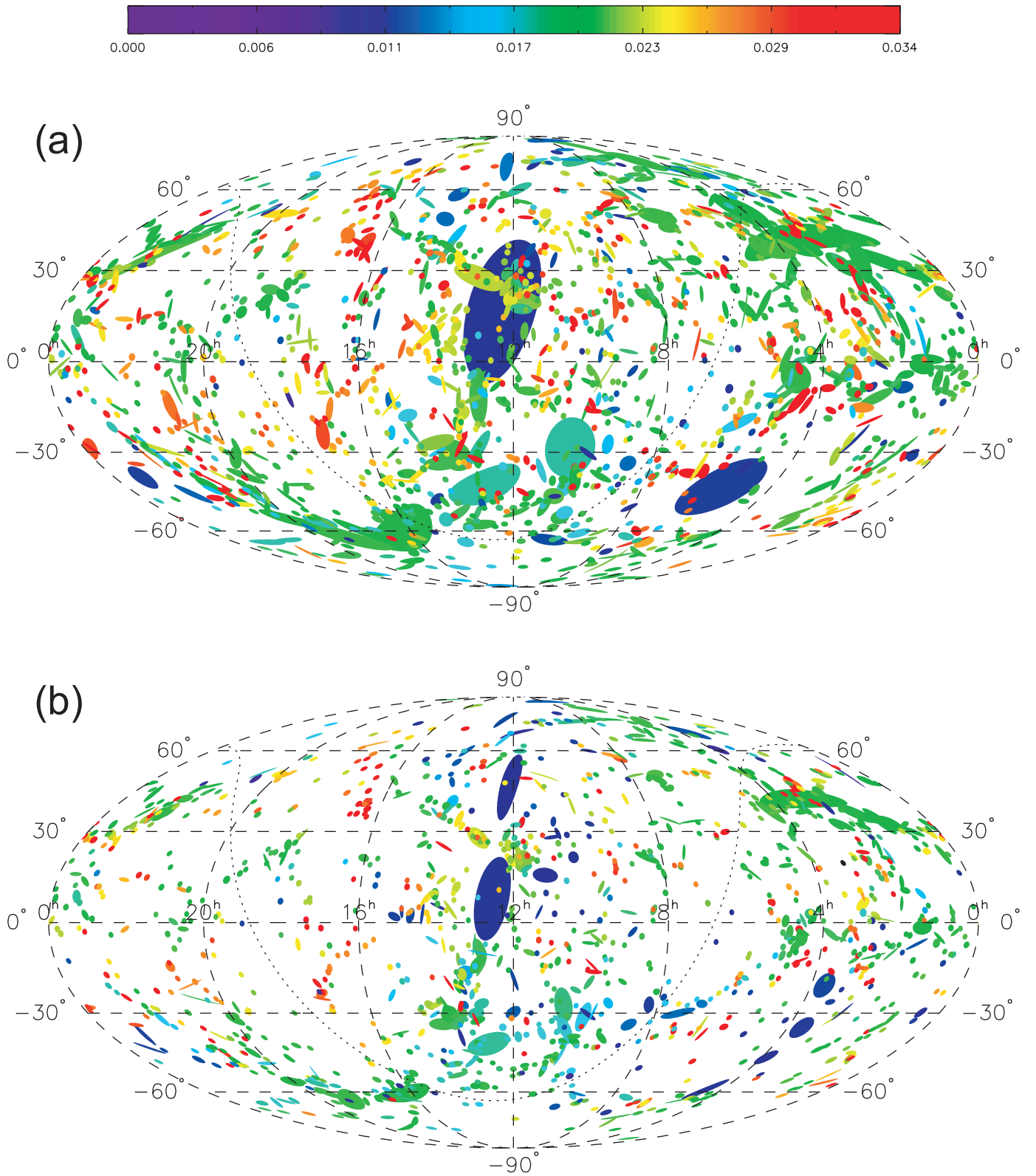


FIG. 5.—Groups (scaled by number of members) identified by clustering algorithm in 2MRS catalog shown in equatorial coordinates. Panels *a* and *b* are shown using a Mollweide projection in equatorial coordinates, centered at $12^{\text{h}}00^{\text{m}}$. Panels *c* and *d* are shown using a Mollweide projection in Galactic coordinates. Panels *a* and *c* show the groups in the LDC catalog ($\delta\rho/\rho = 12$). Panels *b* and *d* show the groups in the HDC catalog ($\delta\rho/\rho = 80$). Each group is plotted as an ellipse with the position angle and axis ratio of the group. The areas of the ellipses are proportional to the number of members in the group, scaled such that the major axis of the largest group is 75% of its true size. The ellipses are transformed from the x - y coordinate system discussed in § 5.3.2 to the appropriate map projection. The color of the ellipse represents the group's mean redshift. The Galactic plane is shown by the dotted line in panels *a* and *b*.

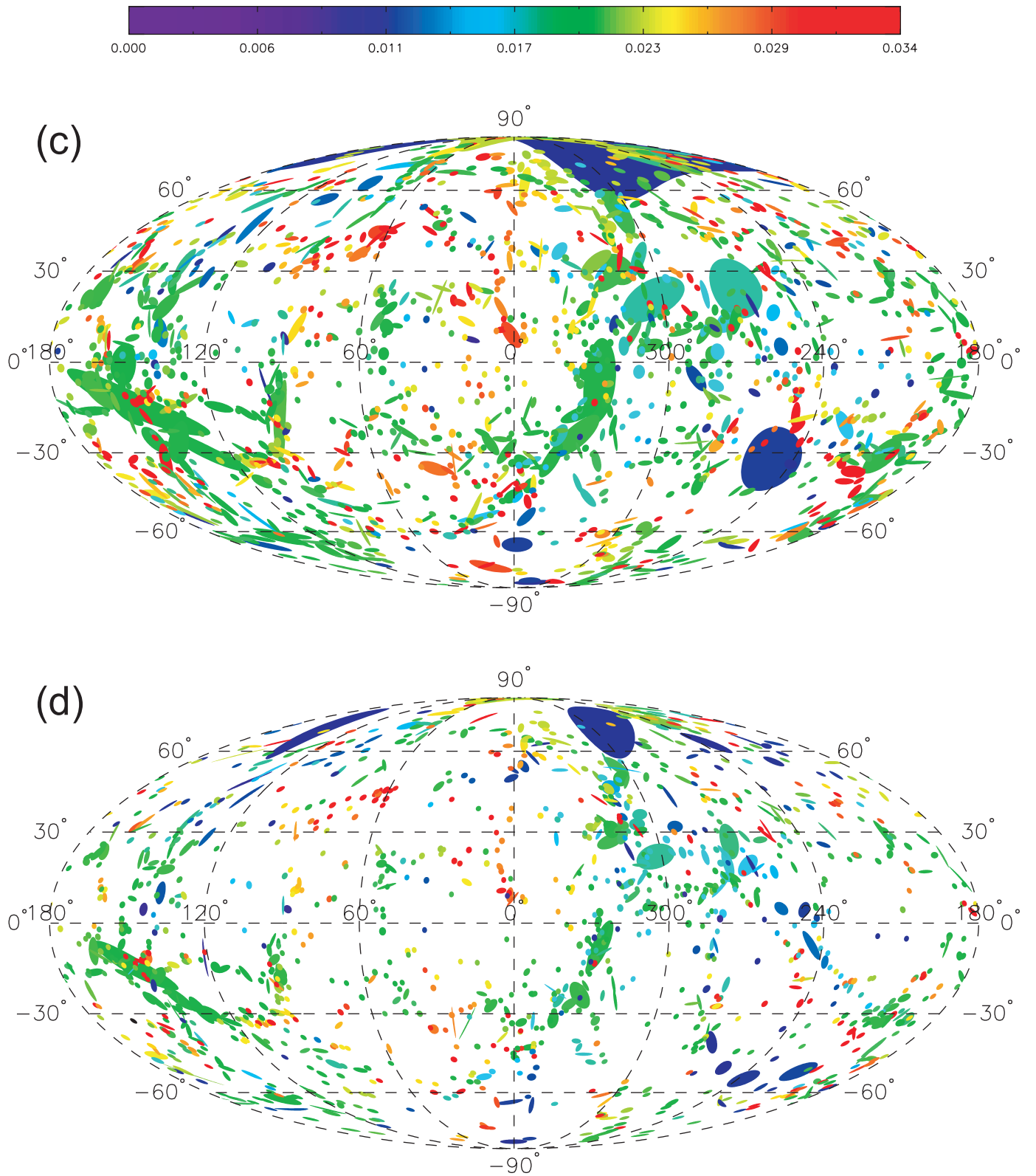


FIG. 5.— *Continued*

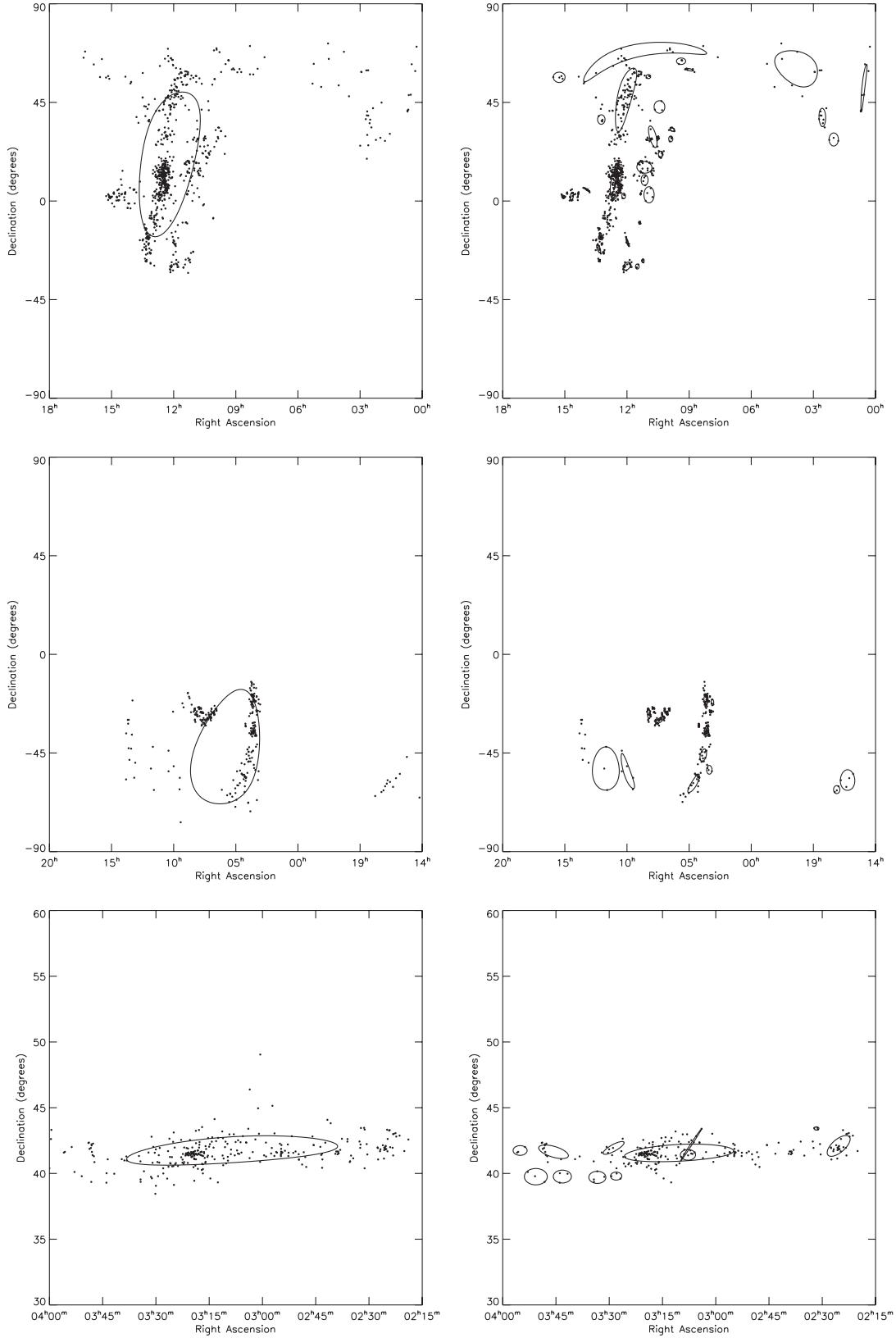


FIG. 6a

FIG. 6.—Six largest groups in the LDC catalog ($\delta p/\rho = 12$). The galaxies in these groups are shown on the plots on the left-hand side. The groups identified in the HDC catalog ($\delta p/\rho = 80$) in the same region of the sky are shown on the right-hand side. The ellipses shown have been computed in the x - y coordinate space discussed in § 5.3.2 and mapped onto the equatorial coordinates used in the figure. The identifications of groups with known clusters and superclusters in (a) are as follows. *Top*: The Virgo, NGC 3607, NGC 4105, IC 764, NGC 5746, NGC 3190, NGC 5846, and NGC 4038 clusters and M81 group (LDC group 852). Note that the ellipse in the right-hand figure at (13 hr, $+10^\circ$) is partially masked by the high density of points. *Middle*: The Eridanus, Fornax I, Dorado, NGC 2280, NGC 1433, and NGC 2559 clusters (LDC group 391). *Bottom*: Perseus-Pisces (A426, A347) (LDC group 229). The identifications in (b) are as follows. *Top*: Hydra (A1060) (LDC group 712). *Middle*: Norma (A3627, the GA) (LDC group 1117). *Bottom*: Centaurus (A3526) (LDC group 881).

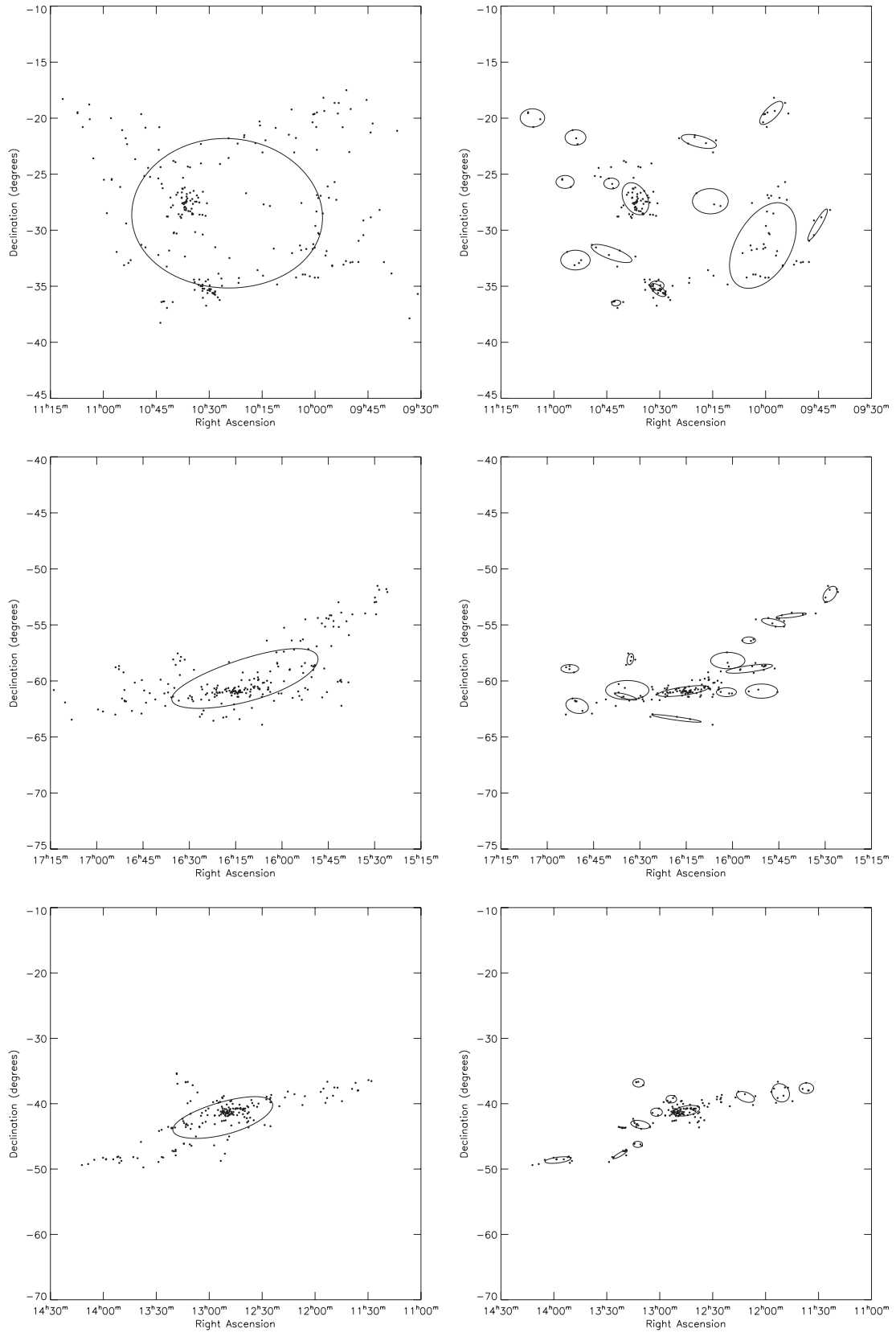


Fig. 6b

TABLE 1
GENERAL PROPERTIES OF THE GROUP CATALOGS

Property	LDC Catalog	HDC Catalog
$\delta\rho/\rho$	12	80
D_0 (Mpc)	1.63	0.89
V_0 (km s ⁻¹).....	399	350
Number of singles.....	5548 (27.2%)	9608 (47.2%)
Number of binaries.....	1397 (13.7%)	1710 (16.8%)
Number of groups of 3+	1538 (59.1%)	1258 (36.1%)
Number of groups of 10+	203 (30.5%)	113 (13.3%)
Number of groups of 50+	17 (10.7%)	8 (4.7%)
Mean number per group	7.84	5.84
Standard deviation	26.25	11.85
Min/max per group.....	3/810	3/298

NOTE.—The values in parentheses represent the percentages of galaxies that fall into this category.

in the LDC catalog contain a higher fraction of interlopers than the groups in the HDC catalog; however, it is evident that the LDC catalog identifies the largest structures on the sky. This suggests that the LDC catalog will be the better candidate for the basis of a flow-field model, as some of the largest structures are fragmented in the HDC catalog. This is explored further in follow-up work.⁹

5.3. Group Properties

In this section, we discuss the properties of the LDC and HDC catalogs. Estimates of the velocity dispersion, size, mass, and luminosity are discussed in § 5.3.1; we obtain estimates of the axis ratio and position angle of the groups in § 5.3.2 below.

5.3.1. Basic Properties

The properties of the LDC and HDC catalogs are summarized in Table 1. In further analysis, we only consider groups with five or more members that are also present in the 2MRS catalog (referred to as *genuine* hereafter), as opposed to those generated in the population of the Galactic plane. We do this in an attempt to exclude groups with a high fraction of interlopers in our analysis.

We provide two estimates of the mass of the groups. We first compute the virial mass of the group, M_V ,

$$M_V = \frac{3\pi}{2} \frac{\sigma_P^2 R_P}{G}, \quad (6)$$

where G is Newton's constant, σ_P is the projected velocity dispersion,

$$\sigma_P^2 = \frac{\sum_i (V_i - V_G)^2}{N - 1}, \quad (7)$$

and R_P is the projected virial radius,

$$R_P = \frac{N(N - 1)}{\sum_{i>j} R_{ij}^{-1}}. \quad (8)$$

Here V_G is the mean group velocity and V_i is the line-of-sight velocity of the i th member. N is the total number of galaxies in the group and R_{ij} is the projected separation between two galaxies, defined in terms of their angular separation θ_{ij} through

$$R_{ij} = \frac{2V_G}{H_0} \tan\left(\frac{\theta_{ij}}{2}\right).$$

⁹ See A. C. Crook et al. (2007b, in preparation).

TABLE 2
MEDIAN PROPERTIES OF GROUPS WITH FIVE OR MORE GENUINE MEMBERS

Property	LDC Catalog	HDC Catalog
σ_P (km s ⁻¹)	197 (183, 206)	183 (166, 193)
R_{PV} (Mpc).....	1.71 (1.58, 1.85)	0.97 (0.90, 1.04)
$\log(M_V/M_\odot)$	13.79 (13.72, 13.90)	13.54 (13.46, 13.60)
$\log(M_P/M_\odot)$	14.05 (13.98, 14.10)	13.66 (13.57, 13.75)
$\log((M_V/L_K)/(M_\odot/L_\odot))$	1.70 (1.63, 1.76)	1.53 (1.45, 1.59)
$\log((M_P/L_K)/(M_\odot/L_\odot))$	1.90 (1.85, 1.97)	1.67 (1.60, 1.73)
$\Omega_{M,V}$	0.14 (0.12, 0.16)	0.10 (0.08, 0.11)
$\Omega_{M,P}$	0.23 (0.20, 0.27)	0.13 (0.11, 0.16)

NOTE.—The median values are shown, with 99% confidence levels in parentheses. We compute the confidence levels by drawing an equally sized sample from the observed distributions and computing the median values 5000 times.

Due to the biases in the virial mass estimator (e.g., Bahcall & Tremaine 1981), we also calculate the projected mass estimator M_P (Bahcall & Tremaine 1981; Heisler et al. 1985),

$$M_P = \frac{f_{PM}}{\pi G(N - \gamma)} \sum_i s_i (V_i - V_G)^2, \quad (9)$$

where s_i is the offset of the i th member (in physical units) from the center of the group. Following Heisler et al. (1985) we set $\gamma = 1.5$ and $f_{PM} = 10.2$.

We estimate the total isophotal K -band luminosity L_K of each cluster from the observed luminosity using the equation

$$L_K = \left[1 - \frac{\gamma(\alpha + 2, L_{\lim}/L^*)}{\Gamma(\alpha + 2)} \right]^{-1} L_{\text{obs}}, \quad (10)$$

where L_{obs} is the total isophotal observed K -band luminosity and L_{\lim} is the limiting observable luminosity at the distance of the cluster d_c ,

$$L_{\lim} = 10^{0.4[M_{\odot,K} - m_{\lim} + 25 + 5 \log(d_c/\text{Mpc})]} L_\odot; \quad (11)$$

α and L^* take the values used quoted in equation (4) above, given

$$L^* = 10^{0.4 \log(M_{\odot,K} - M^*)} L_\odot;$$

$\gamma(m, x)$ is the lower incomplete gamma function,

$$\gamma(m, x) = \int_0^x t^{m-1} e^{-t} dt,$$

and

$$\Gamma(m) = \gamma(m, \infty).$$

Using the values of M_V , $V - K$, and $K - K_s$ obtained from Cox (2000), we set the K_s -band magnitude zero point, $M_{\odot,K} = 3.29$. Note that we have not applied an isophotal correction to the luminosities; thus the luminosities presented in this paper are lower than the total K -band luminosities.

Table 2 contains the median properties of the groups in the catalog that have at least five genuine members. Figure 7 shows the fraction of groups containing at least five genuine members as a function of velocity dispersion, projected virial radius, mass, and mass-to-light ratio (using both virial and projected mass estimates).

As expected, the distribution of virial radii widens and increases with the larger choice of D_0 . Similarly, the velocity dispersions

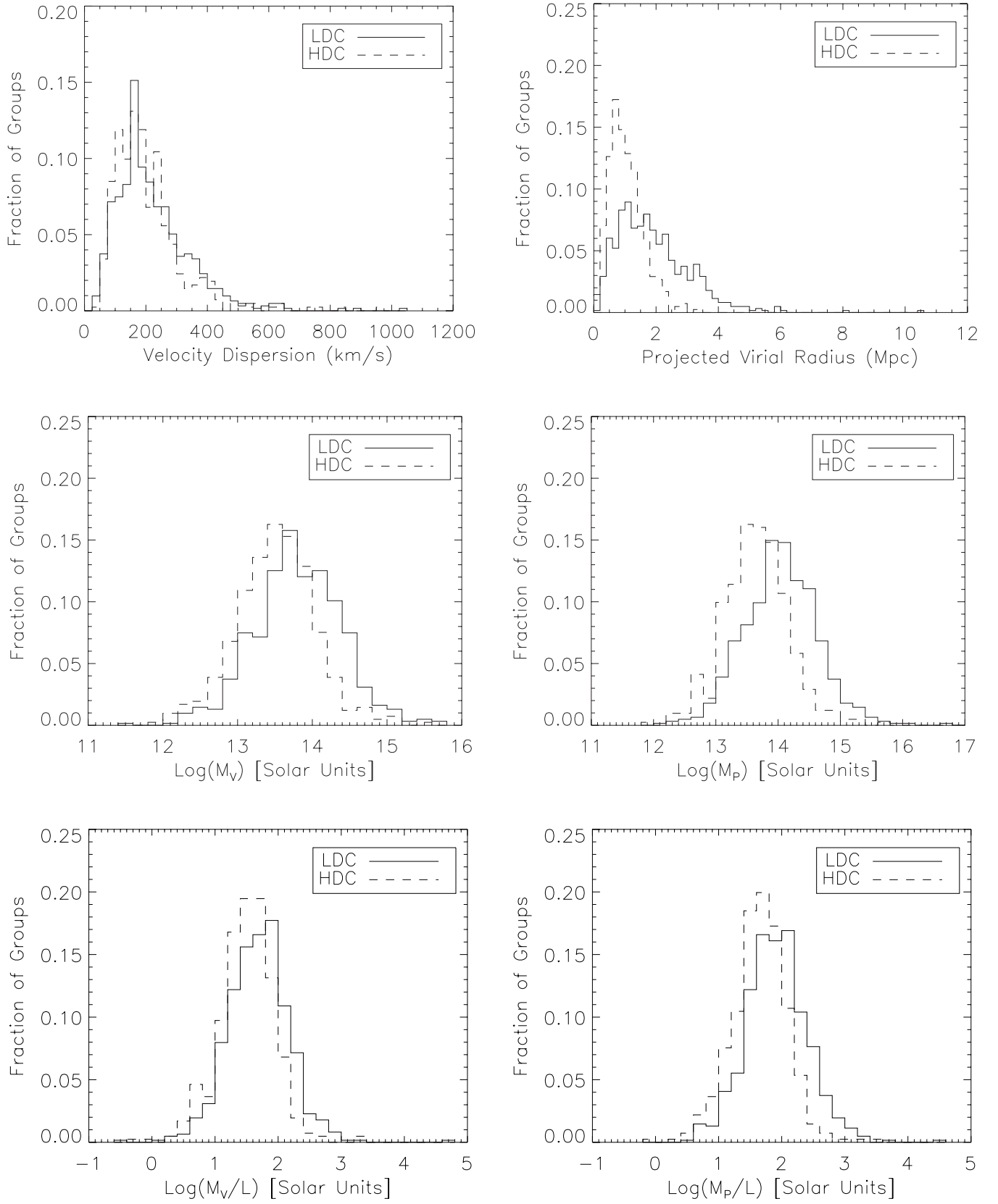


FIG. 7.—Properties of groups with at least five genuine members. We contrast the properties of the groups in the LDC catalog (solid line) with those in the HDC catalog (dashed line). *Top left:* Variation in line-of-sight velocity dispersion across the groups. *Top right:* Projected virial radius of each of the groups. *Middle:* Virial (left) and projected (right) mass estimates of the groups. *Bottom:* Mass-to-light ratios of the groups using the virial (left) and projected (right) mass estimators. We assume $h = 0.73$.

increase with the larger choice of D_0 and V_0 due principally to the change in V_0 ; recalculating the groups for various values of D_0 while keeping V_0 constant has little impact on the distribution of velocity dispersions. The net effect is expected to increase the estimated virial mass of the groups, as observed.

We use the median mass-to-light ratio to obtain a value of Ω_m , the ratio of the matter density ρ_m (assuming all the mass is contained within galaxy clusters) to the critical density ρ_c ,

$$\rho_c = \frac{3H_0^2}{8\pi G}.$$

We use the K -band luminosity density \mathcal{L}_K ,

$$\mathcal{L}_K = \Phi^* L^* \Gamma(\alpha + 2), \quad (12)$$

where α , L^* , and Φ^* take the values given in equation (4), in conjunction with the median mass-to-light ratios to estimate ρ_m ,

$$\rho_{m,E} = \left\langle \frac{M_E}{L_K} \right\rangle \mathcal{L}_K, \quad (13)$$

where E refers to the method of mass estimation (i.e., virial or projected). For the luminosity function and magnitude zero point used in this paper, the critical density corresponds to $M_E/L_K = 353 M_\odot/L_\odot$. Similar results were found by Kochanek et al. (2001) and Bell et al. (2003), adjusting for the authors' isophotal corrections, and magnitude zero points where appropriate. The obtained values are given in Table 2. Spergel et al. (2006) obtain $\Omega_m = 0.238^{+0.013}_{-0.024}$ assuming $h = 0.73$; this value only agrees, at the 1σ level, with the value we obtained using the projected mass estimator in the LDC catalog ($\Omega_m = 0.229^{+0.016}_{-0.012}$, 1σ errors). The virial mass estimates predict Ω_m to be too small. For this reason, we will use the projected mass estimator as opposed to the virial mass estimator in further analysis. The fact that the HDC catalog predicts a value of Ω_m significantly smaller than that obtained using *WMAP* suggests that we are missing a significant fraction of the mass of the cluster in our estimate; selecting groups based on the density contrast $\delta\rho/\rho = 80$ is causing us to underestimate the median mass-to-light ratio of groups. This suggests that the dark matter halos extend beyond the $\delta\rho/\rho = 80$ density-contrast contour that was inferred from luminous matter.

5.3.2. Orientation and Ellipticity

We include in the catalog a measure of the axis ratio, position angle, and semimajor axis of the groups containing five or more genuine members, calculated using the following method: We rotate the coordinates such that center of the group¹⁰ lies along the z -axis. We measure the angle of each galaxy from the z -axis, Θ , as well as its azimuthal angle ψ , then define

$$x = \Theta \cos(\psi), \quad (14)$$

$$y = \Theta \sin(\psi), \quad (15)$$

such that the positive y -axis points north and the positive x -axis points east.¹¹ This definition is chosen such that the shape of the group is not distorted under the projection onto a plane.

We rotate the axes to some angle ϕ and define the coordinates of the galaxies in the rotated frame as $(\tilde{x}_i, \tilde{y}_i)$. We choose the value

¹⁰ We define the center as the mean position of the galaxies in Cartesian coordinates, assuming the galaxies lie on the surface of a unit sphere.

¹¹ Note that the mean values of x and y are not strictly zero but, in the analyzed data, are sufficiently small that an iterative centering procedure is not required.

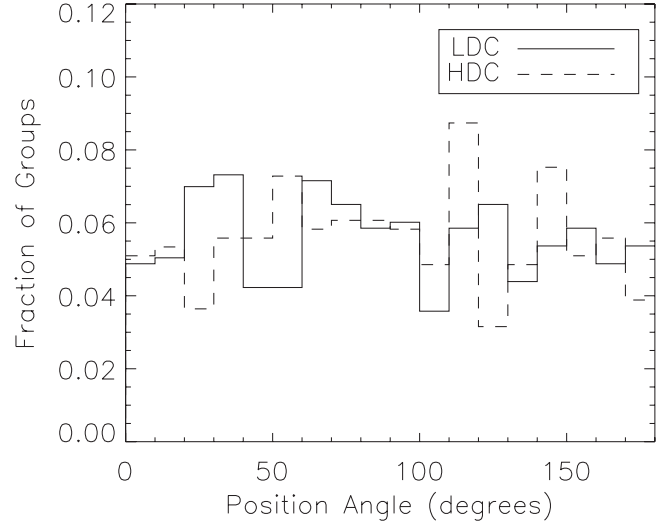


FIG. 8.—Position angles of groups with five or more genuine members. The figure shows the fraction of groups with position angles in the specified intervals (binned by 10°). The groups in the LDC catalog are represented by the solid line, and the HDC catalog by the dashed line.

of ϕ that minimizes $\sum_i \tilde{y}_i^2$. We compute the 75th percentile values of $|\tilde{x}|$ and $|\tilde{y}|$ (\tilde{x}_{75} and \tilde{y}_{75}) and record their ratio η ($0 < \eta < 1$) as a measure of the axis ratio of the group. We also record the larger of \tilde{x}_{75} and \tilde{y}_{75} as a measure of the semimajor axis of the cluster, a , as well as the angle of rotation of the semimajor axis from north toward east (the position angle of the group, ϕ). We verify that these angles are approximately uniformly distributed by showing the number of galaxies as a function of position angle in Figure 8.

The above properties of the groups are illustrated graphically by using ellipses with the same semimajor axes, axis ratios, and position angles in Figure 4 and with the same axis ratios and position angles in Figure 5. For groups with three to four members, a circle is drawn with an angular radius equal to the 75th percentile mean offset.¹²

In Figure 6 we show the galaxies that are associated by the clustering algorithm to form the six largest groups in the LDC catalog. The corresponding ellipses have been overlaid on this plot.¹³ The figure also shows the corresponding groups identified in the HDC catalog. It is evident that the higher density contrast used in the latter choice of parameters splits the large structures identified when choosing a lower density contrast.

5.4. Reliability of the Algorithm

In this section we discuss the verifications performed to ensure that the groups obtained are consistent with both expectation and the literature. In § 5.4.1 we examine the distance dependence of the velocity dispersions and mass-to-light ratios of the groups. We compute the mass functions of the group catalogs and compare them with expectation in § 5.4.2. Finally, we compare the 2MRS group catalogs directly with the UZC-SSRS2 and CfAN group catalogs in § 5.4.3.

5.4.1. Variation with Distance

Figure 9 shows the velocity dispersion σ_P of the groups as a function of distance. We fit a curve of the form

$$\sigma_P = 10^{(\alpha D + \beta)} \text{ km s}^{-1} \quad (16)$$

¹² This value is also reported under the column titled a in Tables 6 and 7.

¹³ Note that the shapes of the ellipses have been distorted due to the choice of coordinate system.

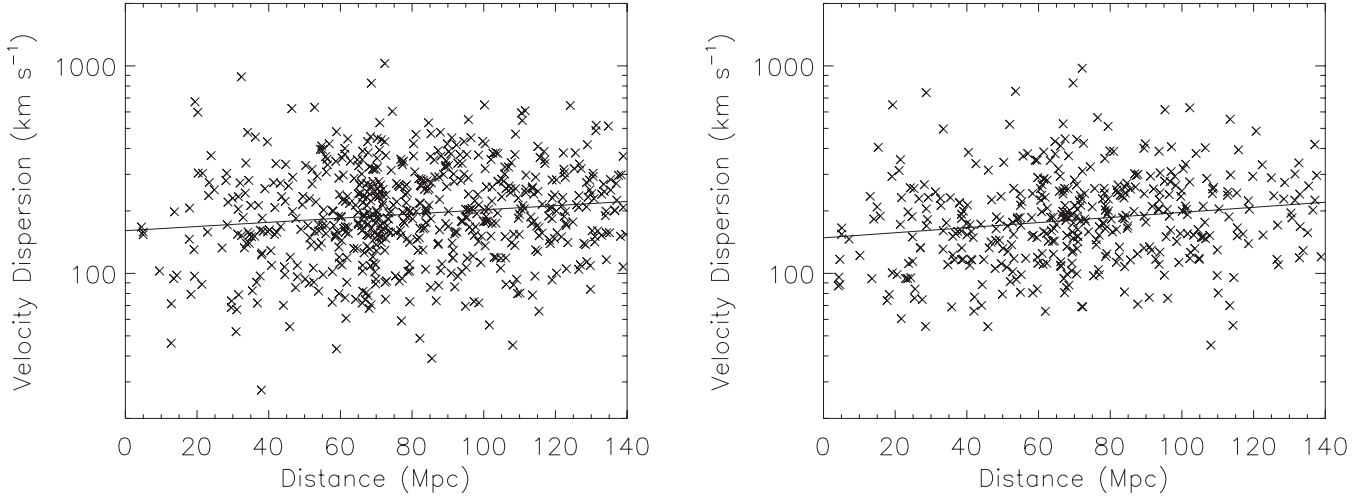


FIG. 9.—Velocity dispersion of groups as a function of distance. The left panel shows the groups in the LDC catalog ($\delta\rho/\rho = 12$). The right panel shows the groups in the HDC catalog ($\delta\rho/\rho = 80$). Only groups with five or more genuine members are included. The solid line shows a linear fit to the data, as described in the text.

to the data, where the best-fit parameters are given in Table 3. The large scatter and very small correlation observed in Figure 9 (the mean velocity dispersion changes by a factor of $0.63 \bar{\sigma}$ between $[20, 40]$ Mpc and $[120, 140]$ Mpc in the HDC catalog, and $0.35 \bar{\sigma}$ in the LDC catalog¹⁴) demonstrates that there is minimal bias introduced in the velocity dispersion of groups with distance; this was desired in the construction of the algorithm (see § 2). Had we chosen to scale V_0 with distance, we would expect the velocity dispersions of the most distant groups to be larger than observed in this figure. Since the correlation is already slightly positive, scaling V_0 would have introduced a more significant bias with distance.

Figure 10 shows the mass-to-light ratios (computed using the projected mass estimators) as a function of distance for both pairs of parameters.

We fit a curve of the form

$$\frac{M_P}{L_K} = 10^{(pD+q)} \frac{M_\odot}{L_\odot} \quad (17)$$

to the data, where the best-fit parameters are given in Table 3. The lower limit to the mass-to-light ratios computed as a function of distance remains approximately constant, while the upper limit decreases with distance, giving rise to the negative slope. Due to the nature of the flux-limited sample, at the largest distances we are not sensitive to (intrinsically) faint objects, and therefore we are preferentially selecting the brightest groups.

¹⁴ Here $\bar{\sigma}$ represents the average standard deviation, weighting the standard deviation of the velocity dispersions in each interval equally.

TABLE 3
VELOCITY DISPERSION–DISTANCE AND M/L -DISTANCE RELATIONS

Parameter	LDC Catalog	HDC Catalog
Eq. (16):		
α (10^{-3} Mpc $^{-1}$).....	1.0 ± 0.3	1.2 ± 0.3
β	2.21 ± 0.03	2.17 ± 0.03
Eq. (17):		
p (10^{-3} Mpc $^{-1}$).....	-7.0 ± 0.6	-5.8 ± 0.6
q	2.48 ± 0.05	2.06 ± 0.05

NOTE.—Values correspond to the parameters in eqs. (16) and (17) that minimize their respective χ^2 statistics.

The scaling of the linking length, D_L , is designed to produce groups with a similar number of members at all distances. Since we are further correcting the luminosity of the groups to account for those galaxies to which the survey was not sensitive, we expect the mean luminosity of groups to increase with distance. As we have already shown that the velocity dispersions of the groups we find (and hence the estimated masses) are comparatively uncorrelated with distance, we would expect that we should miss those groups with high mass-to-light ratios at the largest distances, as indeed we observe in Figure 10. To correct for this effect, one may introduce a scaling in the linking length in velocity space, V_L , with distance. This would increase the estimated group mass with distance; however, such mass estimates would be based on groups containing many interlopers and thus not accurately represent the mass of the group. Such shortcomings of the percolation algorithm will be discussed in more detail in follow-up work.¹⁵

5.4.2. Mass Functions

The large number of groups in the sample allows us to obtain an accurate estimate of the mass function for groups in the LDC and HDC catalogs. In this section we only consider groups with at least five genuine members at a distance of $10 h^{-1}$ Mpc or greater. We compute the mass function using the $1/V_{\max}$ procedure (e.g., Martínez et al. 2002), whereby each group is weighted by the inverse of the maximum comoving volume $V_{\max}(L_i)$ in which the group remains observable given the flux limit of the survey. L_i is the luminosity of the fifth-brightest member of the group. The differential mass function can then be computed as

$$n(M) = \sum_{|M_i - M| \leq \Delta M} [V_{\max}(L_i)]^{-1}, \quad (18)$$

where M_i are the group masses and ΔM is the (variable) bin width. The results are shown in Figure 11.

We consider analytical differential mass functions of the form suggested by Sheth & Tormen (1999) (also see Jenkins et al. 2001),

$$n(M) = A \frac{\gamma \bar{\rho}}{M^2} \sqrt{\frac{a}{\pi}} \left(\frac{M}{M_0} \right)^{\gamma/2} \times \left[1 + \frac{1}{(2a)^p} \left(\frac{M}{M_0} \right)^{-\gamma p} \right] \exp \left(- \left(\frac{M}{M_0} \right)^\gamma \right), \quad (19)$$

¹⁵ See A. C. Crook et al. (2007a, in preparation).

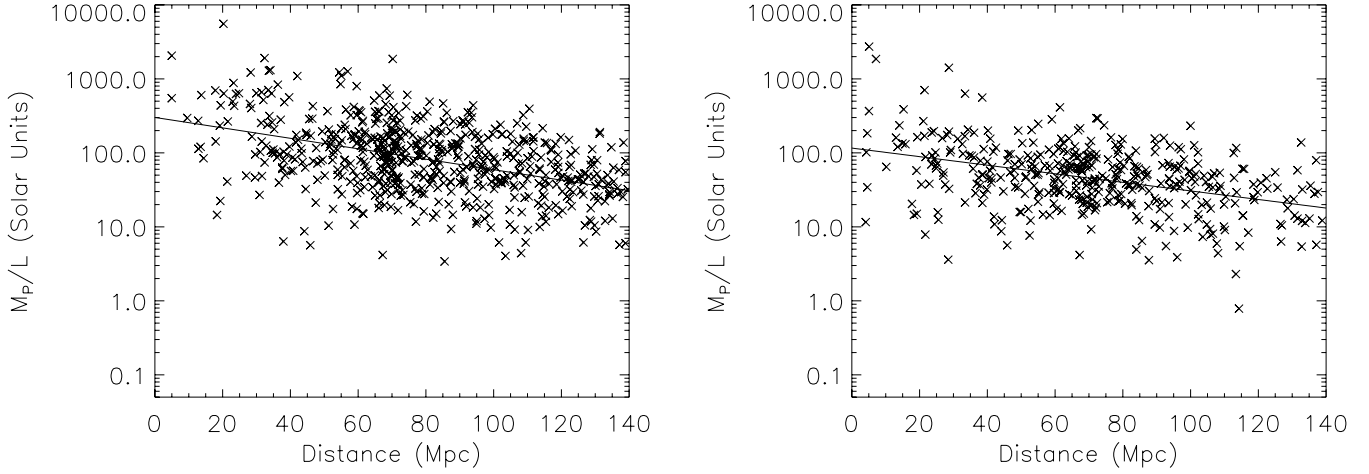


FIG. 10.—Mass-to-light ratio of groups as a function of distance. The plots show the ratio of the projected mass estimate to the corrected K -band luminosity of the cluster. The left panel shows the groups in the LDC catalog ($\delta\rho/\rho = 12$). The right panel shows the groups in the HDC catalog ($\delta\rho/\rho = 80$). Only groups with five or more genuine members are included. The solid line shows a linear fit to the data, as described in the text.

where $\gamma = 1 + (\tilde{n}/3)$ and we set $\tilde{n} = 1$. The choice of parameters $a = 1, p = 0, A = 0.5$ corresponds to the analytical prediction of Press & Schechter (1974, hereafter PS74). Sheth & Tormen (1999, hereafter ST99) suggest the alternative choice parameters $A = 0.3222, a = 0.707, p = 0.3$, which provide good agreement with a subset of N -body simulations analyzed by Jenkins et al. (2001).

We fit functions of the form of equation (19) to groups in the LDC and HDC catalogs using the parameter choices of both PS74 and ST99 (see Fig. 11 and Table 4). The analytical descriptions are both good approximations to the data; similar conclusions were also drawn by Martínez et al. (2002); however, we find the fit to the PS74 form produces a slightly smaller χ^2 statistic in both cases.

We compare the ratio of the best-fit values for M_0 in the PS74 form of the mass function to the value predicted using the simple

arguments of PS74. M_0 scales with the minimum density contrast according to

$$M_0 \propto \left(\frac{\delta\rho}{\rho} \right)^{-2/\gamma}, \quad (20)$$

and thus we expect the ratio of the determined values of M_0 to be given by

$$\begin{aligned} \log \left(\frac{M_0^{(\text{LDC})}}{M_0^{(\text{HDC})}} \right) &= \log \left(\left[\frac{(\delta\rho/\rho)_{\text{LDC}}}{(\delta\rho/\rho)_{\text{HDC}}} \right]^{-2/\gamma} \right) \\ &= -\frac{2}{\gamma} \log \left(\frac{12}{80} \right) = 1.24. \end{aligned}$$

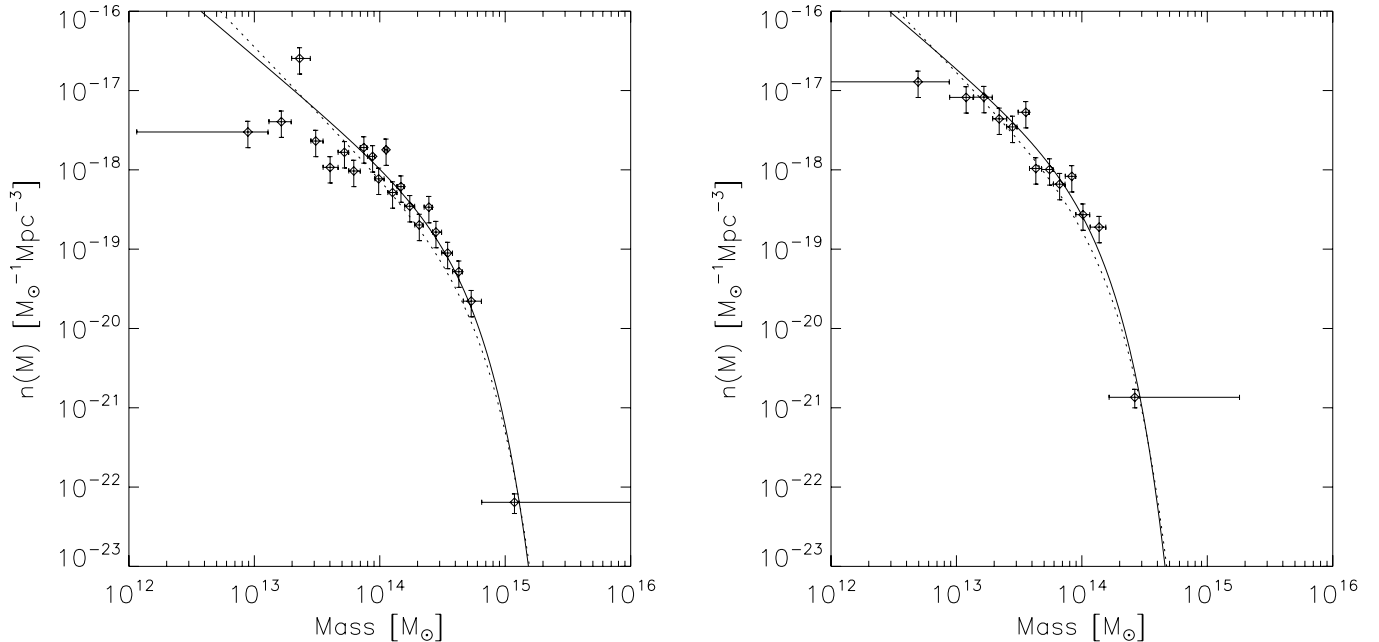


FIG. 11.—Mass functions of the groups. The left panel shows the differential mass function of the groups in the LDC catalog ($\delta\rho/\rho = 12$); the right panel shows that of the groups in the HDC catalog ($\delta\rho/\rho = 80$). The masses are estimated using the projected mass estimator; 2σ error bars are shown on the plots. Only groups at a distance of at least 10 Mpc containing at least five genuine members are included. The points have been fitted with a Press-Schechter (PS74) mass function (solid line) and a Sheth-Tormen (ST99) mass function (dotted line). The former fit produces the smaller χ^2 statistic in both cases. We assume $h = 0.73$.

TABLE 4
BEST-FIT VALUES FOR PRESS-SCHECHTER AND SHETH-TORMEN MASS FUNCTIONS

Parameter	LDC Catalog	HDC Catalog
PS74 Form ^a		
$\bar{\rho}$ ($10^{10} M_{\odot} \text{ Mpc}^{-3}$).....	2.8 ± 1.4	0.9 ± 0.6
$\log(M_0/M_{\odot})$	14.50 ± 0.13	13.93 ± 0.18
ST99 Form ^b		
$\bar{\rho}$ ($10^{10} M_{\odot} \text{ Mpc}^{-3}$).....	3.6 ± 1.9	1.3 ± 0.8
$\log(M_0/M_{\odot})$	14.43 ± 0.16	13.87 ± 0.21

NOTE.—Values correspond to the parameters in eq. (19) that minimize the χ^2 statistic.

^a Press & Schechter (1974).

^b Sheth & Tormen (1999).

The obtained values of M_0 give a log ratio of 0.57 ± 0.22 , providing agreement only at the 3σ level with Press-Schechter theory. The two methods cannot be expected to be in perfect agreement as the group-identification algorithm will find different groups as V_0 is varied but D_0 is held constant. The computed estimates of the group masses will therefore vary, while the ratio predicted by the Press-Schechter treatment (which is sensitive only to the change in density contrast, thus D_0) does not.

5.4.3. Comparison with Other Group Catalogs

To verify the validity of the produced group catalogs, we compare the 2MASS group catalogs with the UZC-SSRS2 group catalog (Ramella et al. 2002) and the CfAN group catalog (Ramella et al. 1997). The former is constructed from partial versions of the Updated Zwicky Catalog (UZC)¹⁶ and the Southern Sky Redshift Survey (SSRS2; da Costa et al. 1998) and covers 37% of the sky; the latter covers 10%. The selection criteria for the three catalogs result in completeness limits at different velocities. To compute the completeness in redshift space, we use the same technique as in § 3.4 above; however, in this case we determine the velocity corresponding to the peak in the derivative of equation (5). When comparing the 2MASS group catalog with the other two, we cut both catalogs at the smallest of the two corresponding velocity limits (2MRS at 5697 km s^{-1} in the LDC catalog and 5350 km s^{-1} in the HDC catalog, UZC at 7115 km s^{-1} , CfAN at 9390 km s^{-1}).

For each of the groups present in the UZC-SSRS2 and CfAN catalogs, we search the 2MRS group catalogs for a group within a radius set by twice the sum of the virial radii of the group in 2MRS and the group in the comparison catalog, with mean velocities that differ by less than 30%. We find that 86% of the groups in UZC-SSRS2 and 76% of the groups in CfAN are present in the 2MRS LDC group catalog, whereas 78% of the groups in UZC-SSRS2 and 69% of the groups in CfAN are present in the 2MRS HDC group catalog. There is reasonable agreement in both cases given the differing selection criteria employed in the three samples.

We also compare the distributions of velocity dispersion using the two-sided Kolmogorov-Smirnov (K-S) test (Chakravarti et al. 1967, p. 392). The cumulative fractions of the LDC and HDC catalogs are compared with the cumulative fractions of the UZC-SSRS2 and CfAN catalogs (see Fig. 12). For comparison, we also show the velocity dispersions of groups in Sloan Digital Sky Survey Data Release 3 (SDSSDR3; Merchán & Zandivarez 2005). The results of the test are shown in Table 5.

The velocity dispersions in both the LDC and HDC catalogs are consistent with the CfAN sample. However, only the LDC catalog produces groups with velocity dispersions consistent with those in the UZC-SSRS2 catalog, although the discrepancy with the HDC catalog is small. Any discrepancy with the CfAN catalog is not appreciable due to the smaller number of groups in the compared sample of CfAN galaxies.¹⁷ Neither the LDC or HDC catalogs are consistent with the SDSSDR3 group catalog. The discrepancy is expected due to the difference in the algorithms used to identify groups.

The difference in the distributions of velocity dispersion is due in part to the different scalings of the velocity linking parameter. When scaling the linking length, the more distant groups are likely to have higher velocity dispersion. These groups may not be present in a group catalog derived by setting the parameter to a constant. The velocity dispersions of groups in the UZC-SSRS2 and CfAN catalogs tail off at $\sim 400 \text{ km s}^{-1}$. The LDC catalog is in reasonable agreement with this since the velocity linking parameter is $\sim 400 \text{ km s}^{-1}$. The HDC catalog, however, contains a smaller fraction of groups with velocity dispersions between 350 and 400 km s^{-1} than are present in the other catalogs. This is consistent with the above hypothesis since, in this case, the velocity linking parameter is set to 350 km s^{-1} . Without knowledge of the complete phase-space positions of the galaxies, we cannot determine

¹⁶ VizieR Online Data Catalog, 611, 10438 (E. E. Falco et al., 1999).

¹⁷ This is accounted for in the K-S test.

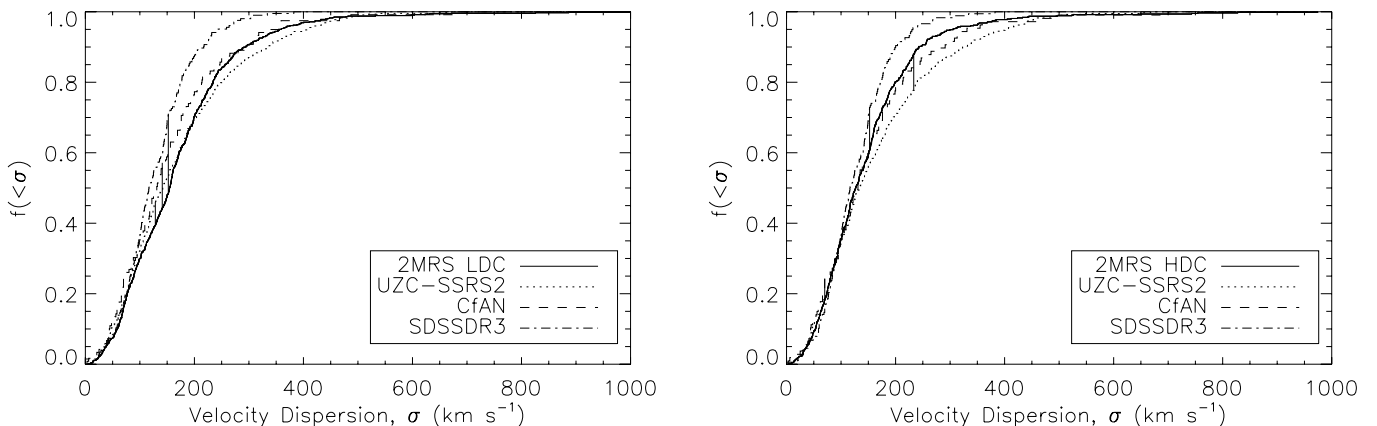


FIG. 12.—Cumulative fraction of velocity dispersions of groups. The velocity dispersions in the UZC-SSRS2, CfAN, and SDSSDR3 group catalogs are compared with the 2MRS LDC catalog (left) and the 2MRS HDC catalog (right). The vertical lines indicate the corresponding D -statistic used in the K-S test.

TABLE 5
COMPARISON BETWEEN 2MRS, UZC-SSRS2, CfAN, AND SDSSDR3 GROUPS

Property	LDC Catalog	HDC Catalog
Number of 2MRS Groups	736	630
Comparison with UZC-SSRS2		
Number of UZC-SSRS2 groups	444	399
D -statistic ^a	0.070	0.102
P -value ^b	0.128	0.011
Comparison with CfAN		
Number of CfAN groups	119	107
D -statistic ^a	0.128	0.062
P -value ^b	0.062	0.862
Comparison with SDSSDR3		
Number of SDSSDR3 groups	204	177
D -statistic ^a	0.227	0.122
P -value ^b	10^{-7}	0.029

NOTE.—The table presents the results of a two-sided Kolmogorov-Smirnov test on the comparison of the distributions of velocity dispersions of the groups in the 2MRS catalog with those in the UZC-SSRS2, CfAN, and SDSSDR3 catalogs.

^a D -statistic is used in computation of P -value in Kolmogorov-Smirnov test.

^b P -value represents the probability that such a difference would be observed under the assumption that the two samples were drawn from the same parent distribution. We consider values of $P < 0.05$ to indicate that the two samples were drawn from significantly different parent distributions.

whether the majority of groups with velocity dispersions greater than 350 km s^{-1} in the LDC, UZC-SSRS2, and CfAN catalogs are bound or not.

We further examine the number density of groups as a function of redshift (see Fig. 13). For comparison, we include the number density of groups in SDSSDR3 and the 2dFGRS Percolation-Inferred Galaxy Group (2PIGG) catalog (Eke et al. 2004), although these surveys have very different selection biases and fainter flux limits, which result in a higher number density of groups than that detected in 2MRS.

The 2MRS group catalogs contain an approximately constant number density of groups to $\sim 80 \text{ Mpc}$, before following a power-law decay. The LDC catalog has slightly more uniform coverage out to this distance than the HDC catalog. The observed shape of the function is consistent between the LDC, HDC, UZC-SSRS2, and CfAN catalogs, and the values only diverge at $\sim 90 \text{ Mpc}$ where the linking length becomes large enough that the selection biases of the parent surveys will have a significant effect on the identified groups. The number density of groups in the SDSSDR3 and 2PIGG catalogs is approximately constant over the entire range of distances considered due to the lower flux limits of the surveys. We note that the HDC catalog contains fewer groups than the LDC catalog beyond $\sim 30 \text{ Mpc}$, as expected. The observed peak at $\sim 70 \text{ Mpc}$ corresponds to the location of the GA and is enhanced by the distance corrections discussed in § 3.1.

6. SUMMARY

We have presented two catalogs of groups in the 2MASS Redshift Survey, identified using a variable-linking-length percolation algorithm (HG82). We discussed the effect of the variation of the input parameters D_0 and V_0 on the number of groups obtained. As demonstrated in § 5.4 above, we see that the correct choice of parameters depends on the purpose of the catalog, and full phase-space information of each galaxy is required to understand the most suitable choice of parameters to find virialized groups.¹⁸

We justify the choice of two pairs of parameters: $(D_0, V_0) = (1.63 \text{ Mpc}, 399 \text{ km s}^{-1})$, corresponding to a density contrast $\delta\rho/\rho = 12$ (LDC catalog), and $(D_0, V_0) = (0.89 \text{ Mpc}, 350 \text{ km s}^{-1})$, corresponding to a density contrast $\delta\rho/\rho = 80$ (HDC catalog). We show that the latter choice of parameters identifies the largest nearby clusters individually, while many of these groups are merged with the former parameter choice.

We compute virial and projected mass estimates for the clusters under the assumptions that the identified groups have spherical symmetry and that the light traces the distribution of the mass. We find that the projected mass estimates give mass functions in

¹⁸ See follow-up work, A. C. Crook et al. (2007a, in preparation).

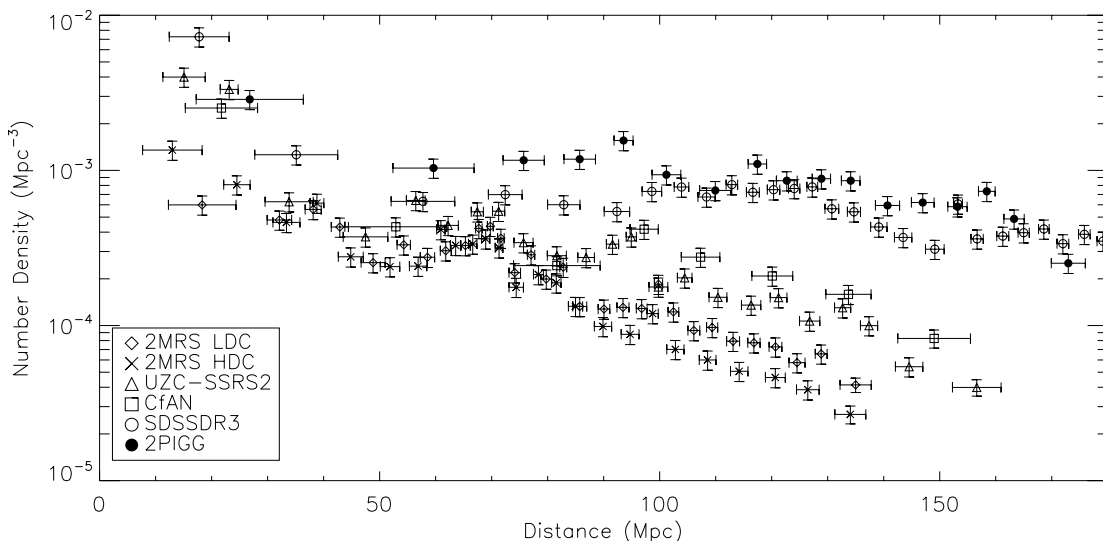


FIG. 13.—Number density of groups as a function of distance. We show the number density for the LDC (diamonds) and HDC (crosses) catalogs, as well as the UZC-SSRS2 (triangles), CfAN (squares), SDSSDR3 (open circles), and 2PIGG (filled circles) catalogs. (The symbols are colored black, blue, green, red, orange, and yellow, respectively, in the electronic edition.) We assume $h = 0.73$. [See the electronic edition of the Journal for a color version of this figure.]

agreement at the 3σ level with Press-Schechter theory, although given the dependence of the algorithm on V_0 , 1σ agreement is not expected.

We calculate corrected K -band luminosities for each cluster, and use these to estimate the mass-to-light ratios and corresponding values of Ω_m . The values predicted using the virial mass estimator in both the LDC and HDC catalogs are significantly smaller than the 3 year *WMAP* result (Spergel et al. 2006) of $\Omega_m = 0.238^{+0.013}_{-0.024}$, motivating the use of the projected mass estimator over the virial mass estimator in subsequent analysis. The projected mass estimates of groups in the LDC catalog produce a value of $\Omega_m = 0.229^{+0.016}_{-0.012}$, which agrees with the *WMAP* result at the 1σ level. The HDC catalog significantly underpredicts the *WMAP* value, suggesting that by only including groups with density contrasts $\delta\rho/\rho \geq 80$ we are underestimating the total mass in groups.

The distribution of velocity dispersions of groups in the 2MRS LDC catalog is in agreement with the groups in the UZC-SSRS2 catalog (Ramella et al. 2002), as well as with those in the CfAN group catalog. The 2MRS HDC catalog velocity dispersions are in agreement with the groups of the CfAN catalog; however, we find that there is a statistically significant difference between the distributions of velocity dispersions in the 2MRS HDC and UZC-SSRS2 catalogs.

The group catalogs presented in this paper provide an estimate of cluster locations and masses without the necessity of assuming an intrinsic mass-to-light ratio. We use and discuss the results of the clustering analysis in follow-up work,¹⁹ including the development of a flow-field model to estimate the discrepancy between the expected flow of the local group and the observed

dipole in the cosmic microwave background (Bennett et al. 1996, 2003). These group catalogs form the basis for a map of baryonic density enhancements in the nearby universe, which can be compared with flow-field maps developed using peculiar-velocity surveys in order to infer the presence and location of dark matter that is not correlated with luminous matter.

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¹⁹ See A. C. Crook et al. (2007b, in preparation).

APPENDIX

GROUP CATALOGS

TABLE 6
LDC CATALOG OF GROUPS IN THE 2MASS REDSHIFT SURVEY

Number (1)	R.A. (2)	Decl. (3)	Members (4)	Distance (Mpc) (5)	V_G (km s ⁻¹) (6)	σ_P (km s ⁻¹) (7)	R_P (Mpc) (8)	$\log\left(\frac{M_V}{M_\odot}\right)$ (9)	$\log\left(\frac{M_P}{M_\odot}\right)$ (10)	$\log\left(\frac{M_P/L}{M_\odot/L_\odot}\right)$ (11)	a (arcmin) (12)	η (13)	ϕ (deg) (14)
1.....	00 00 09.5	+32 44 28	3 (0)	136.62	10087	284.5	0.45	13.597	13.398	1.161	5
2.....	00 00 48.3	+04 05 18	3 (0)	119.62	8870	272.9	3.29	14.429	14.320	2.373	39
3.....	00 00 50.1	+28 17 00	5 (0)	120.31	8894	199.7	0.93	13.609	13.930	1.585	33	0.22	80
4.....	00 02 07.0	+06 57 55	4 (0)	71.53	5306	74.7	4.66	13.455	13.505	1.815	122
5.....	00 02 28.8	-54 29 49	3 (0)	132.62	9652	198.6	0.02	12.002	14.117	1.739	33
6.....	00 05 31.5	+27 29 37	3 (0)	102.61	7590	86.8	1.95	13.207	13.206	1.498	29
7.....	00 05 41.8	+05 09 11	3 (0)	71.94	5340	41.0	0.64	12.072	11.974	0.594	15
8.....	00 06 25.3	-52 12 28	3 (0)	138.95	10130	289.9	4.81	14.646	14.715	2.408	51
9.....	00 06 32.0	+32 25 02	10 (0)	66.21	4871	102.6	0.84	12.988	13.406	1.426	86	0.19	40
10.....	00 06 51.6	-33 40 17	3 (0)	93.51	6872	90.5	4.07	13.562	13.523	1.620	60

NOTES.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds. This catalog has been produced using parameters (D_0 , V_0) = (1.63 Mpc, 399 km s⁻¹), corresponding to the density contrast $\delta\rho/\rho = 12$. We assume $h = 0.73$ where a value is required. Col. (4): Number of group members (including those generated from the population of the plane). The number derived from the Galactic-plane population is contained in parentheses. Col. (5): Mean (corrected) group distance. Col. (6): Mean heliocentric group velocity. Col. (7): Line-of-sight velocity dispersion. Col. (8): Projected virial radius. Col. (9): Log of the virial mass in solar units. Col. (10): Log of the projected mass in solar units. Col. (11): Log of the (projected) mass-to-light ratio in solar units. Col. (12): Semimajor axis of the ellipse fit to the group at the 75th percentile level. Col. (13): Axis ratio of ellipse fit to the group members. Col. (14): Position angle of semimajor axis of ellipse fit to the group members; measured from north toward east. Table 6 is published in its entirety in the electronic edition of the *Astrophysical Journal*. A portion is shown here for guidance regarding its form and content.

TABLE 7
HDC CATALOG OF GROUPS IN THE 2MASS REDSHIFT SURVEY

Number (1)	R.A. (2)	Decl. (3)	Members (4)	Distance (Mpc) (5)	V_G (km s ⁻¹) (6)	σ_P (km s ⁻¹) (7)	R_P (Mpc) (8)	$\log\left(\frac{M_V}{M_\odot}\right)$ (9)	$\log\left(\frac{M_P}{M_\odot}\right)$ (10)	$\log\left(\frac{M_P/L}{M_\odot/L_\odot}\right)$ (11)	a (arcmin) (12)	η (deg) (13)	ϕ (deg) (14)	Corresponding Group Number (15)
1.....	00 00 09.5	+32 44 28	3 (0)	136.62	10087	284.5	0.45	13.597	13.398	1.161	5	1
2.....	00 00 39.2	+47 05 02	10 (0)	70.88	5165	151.6	1.60	13.606	13.696	1.652	69	0.32	97	11
3.....	00 02 37.5	+31 20 47	4 (0)	66.16	4867	99.6	0.17	12.278	13.006	1.586	24	9
4.....	00 05 31.5	+27 29 37	3 (0)	102.61	7590	86.8	1.95	13.207	13.206	1.498	29	6
5.....	00 05 41.8	+05 09 11	3 (0)	71.94	5340	41.0	0.64	12.072	11.974	0.594	15	7
6.....	00 09 11.3	+33 07 36	6 (0)	66.25	4873	113.9	1.05	13.174	13.328	1.488	23	0.69	46	9
7.....	00 10 35.7	-56 59 21	3 (0)	132.31	9620	177.4	0.51	13.242	13.365	1.016	7	13
8.....	00 12 00.7	+16 09 37	3 (0)	13.51	912	101.3	1.18	13.123	13.216	2.390	174	16
9.....	00 13 05.5	+30 57 12	3 (0)	65.04	4792	61.3	0.56	12.360	12.424	0.969	12	15
10.....	00 13 18.8	+22 17 48	3 (0)	81.42	6034	121.2	0.58	12.973	13.454	1.879	31	14

NOTES.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds. This catalog has been produced using parameters $(D_0, V_0) = (0.89 \text{ Mpc}, 350 \text{ km s}^{-1})$, corresponding to the density contrast $\delta\rho/\rho = 80$. We assume $h = 0.73$ where a value is required. Col. (4): Number of group members (including those generated from the population of the plane). The number derived from the Galactic-plane population is contained in parentheses. Col. (5): Mean (corrected) group distance. Col. (6): Mean heliocentric group velocity. Col. (7): Line-of-sight velocity dispersion. Col. (8): Projected virial radius. Col. (9): Log of the virial mass in solar units. Col. (10): Log of the projected mass in solar units. Col. (11): Log of the (projected) mass-to-light ratio in solar units. Col. (12): Semimajor axis of the ellipse fit to the group at the 75th percentile level. Col. (13): Axis ratio of ellipse fit to the group members. Col. (14): Position angle of semimajor axis of ellipse fit to the group members; measured from north toward east. Col. (15): Group number from LDC catalog that encompasses all members of this group. Table 7 is published in its entirety in the electronic edition of the *Astrophysical Journal*. A portion is shown here for guidance regarding its form and content.

TABLE 8
GROUPS IN THE LDC CATALOG WITH 50 OR MORE MEMBERS

Number (1)	R.A. (2)	Decl. (3)	Members (4)	Distance (Mpc) (5)	V_G (km s ⁻¹) (6)	σ_P (km s ⁻¹) (7)	R_P (Mpc) (8)	$\log\left(\frac{M_V}{M_\odot}\right)$ (9)	$\log\left(\frac{M_P}{M_\odot}\right)$ (10)	$\log\left(\frac{M_P/L}{M_\odot/L_\odot}\right)$ (11)	Identified with (12)
852.....	12 19 12.0	+16 56 52	810 (2)	19.46	1314	671.4	10.42	15.712	16.102	2.641	Virgo + NGC 3607 cluster
391.....	05 29 43.7	-42 50 28	302 (22)	20.25	1493	597.8	8.07	15.500	16.728	3.745	Eridanus + Fornax I + NGC 2280 cluster + Dorado
229.....	03 08 45.7	+41 47 55	301 (0)	72.39	5322	1027.5	4.95	15.758	15.982	2.435	Perseus-Pisces (A426, A347)
712.....	10 24 57.5	-28 29 04	241 (0)	46.47	3283	623.4	5.96	15.405	15.681	2.463	Hydra (A1060, A1060)
1117.....	16 10 44.6	-59 50 30	217 (42)	68.66	4815	825.7	4.35	15.512	15.723	2.261	Norma (A3627)
881.....	12 51 27.6	-42 11 34	202 (0)	52.79	3301	631.3	4.65	15.308	15.347	1.976	Centaurus (A3526)
1259.....	18 57 14.6	-62 19 13	103 (0)	62.79	4415	348.3	3.70	14.691	14.871	1.771	
79.....	01 15 17.5	+32 47 01	98 (0)	67.60	5006	431.7	3.35	14.836	15.071	1.974	
280.....	04 03 29.2	+51 34 16	86 (78)	71.02	5182	532.0	6.10	15.277	15.303	2.283	
956.....	13 39 37.8	-31 05 29	84 (0)	58.92	4322	481.8	3.25	14.917	15.204	2.237	Centaurus (A3574)
890.....	12 58 53.8	+27 53 23	84 (0)	100.24	6842	648.0	3.31	15.183	15.271	1.916	Coma (A1656)
883.....	12 54 53.4	-11 16 36	82 (0)	62.08	4258	444.7	3.37	14.863	15.186	2.255	
328.....	04 41 15.9	-05 08 37	77 (0)	61.24	4510	375.3	4.12	14.803	15.040	2.177	
1454.....	22 28 34.8	+35 40 15	76 (0)	84.62	6143	532.0	4.23	15.118	15.445	2.295	
127.....	01 57 18.5	+34 19 08	71 (0)	66.22	4905	388.1	3.05	14.702	14.966	2.043	Perseus-Pisces (A262)
811.....	11 45 24.1	+20 08 13	56 (0)	95.74	6548	550.8	1.83	14.785	14.825	1.731	Coma (A1367)
292.....	04 14 23.7	+36 58 31	51 (0)	81.93	6017	286.1	3.49	14.496	14.810	1.867	

NOTES.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds. The LDC group catalog was created using $(D_0, V_0) = (1.63 \text{ Mpc}, 399 \text{ km s}^{-1})$, corresponding to the density contrast $\delta\rho/\rho = 12$. We assume $h = 0.73$ where a value is required. Col. (4): Number of group members (including those generated from the population of the plane). The number derived from the Galactic-plane population is contained in parentheses. Col. (5): Mean (corrected) group distance. Col. (6): Mean heliocentric group velocity. Col. (7): Line-of-sight velocity dispersion. Col. (8): Projected virial radius. Col. (9): Log of the virial mass in solar units. Col. (10): Log of the projected mass in solar units. Col. (11): Log of the (projected) mass-to-light ratio in solar units. Col. (12): Composition of group based on known galaxy clusters and superclusters (lists only those groups that appear in HDC catalog with 25 or more members).

TABLE 9
GROUPS IN THE HDC CATALOG WITH 25 OR MORE MEMBERS

Number (1)	R.A. (2)	Decl. (3)	Members (4)	Distance (Mpc) (5)	V_G (km s ⁻¹) (6)	σ_P (km s ⁻¹) (7)	R_P (Mpc) (8)	$\log\left(\frac{M_V}{M_\odot}\right)$ (9)	$\log\left(\frac{M_P}{M_\odot}\right)$ (10)	$\log\left(\frac{M_P/L}{M_\odot/L_\odot}\right)$ (11)	Corresponding Group Number (12)	Identified with (13)
716.....	12 33 06.5	+07 47 52	298 (0)	19.33	1353	648.7	2.95	15.134	15.257	2.212	852	Virgo
214.....	03 10 24.3	+41 34 04	172 (0)	72.21	5310	974.2	2.61	15.434	15.769	2.470	229	Perseus-Pisces (A426)
699.....	12 08 34.1	+46 13 02	123 (0)	15.30	966	405.0	3.39	14.785	15.109	2.589	852	Virgo
722.....	12 44 48.4	-41 05 22	100 (0)	53.67	3452	752.1	1.95	15.083	15.193	2.137	881	Centaurus (A3526)
928.....	16 15 53.1	-60 54 29	90 (0)	69.67	4887	827.4	1.46	15.039	14.997	1.892	1117	Norma (A3627)
592.....	10 36 54.5	-27 11 21	67 (0)	51.97	3654	523.0	1.48	14.648	14.637	1.967	712	Hydra (A1060)
1034.....	18 47 40.8	-63 26 02	51 (0)	63.29	4450	350.1	1.75	14.372	14.476	1.706	1259	
237.....	03 38 20.8	-20 34 53	51 (0)	21.77	1624	231.8	1.79	14.022	14.106	1.946	391	Eridanus
235.....	03 35 11.8	-35 04 15	43 (0)	19.11	1438	293.2	1.02	13.981	14.094	1.761	391	Fornax I
734.....	12 58 51.0	+27 51 00	42 (0)	102.23	6986	628.2	1.70	14.866	14.955	1.855	890	Coma (A1656)
669.....	11 44 35.4	+19 58 45	42 (0)	95.31	6517	614.5	1.22	14.701	14.775	1.813	811	Coma (A1367)
553.....	10 00 36.8	-31 21 56	40 (0)	37.97	2706	218.4	2.67	14.144	14.371	1.894	712	Hydra (A1060)
628.....	11 09 31.5	+15 23 42	39 (0)	13.04	1035	235.2	1.16	13.845	14.082	2.205	852	NGC 3607 cluster
103.....	01 53 36.6	+36 17 59	39 (0)	66.12	4891	412.0	1.55	14.461	14.466	1.810	127	Perseus-Pisces (A262)
828.....	14 01 45.9	-33 50 04	36 (0)	55.71	4203	367.7	1.59	14.371	14.419	1.975	985	
73.....	01 23 18.6	+33 34 59	35 (0)	66.89	4953	526.0	1.01	14.486	14.429	1.806	79	
584.....	10 30 25.5	-35 20 24	30 (0)	40.47	2928	383.0	0.69	14.047	13.935	1.660	712	
727.....	12 52 38.0	-08 57 53	29 (0)	59.59	4070	350.8	1.34	14.258	14.367	1.984	883	
141.....	02 25 24.5	+42 05 38	29 (0)	76.52	5642	563.1	1.36	14.675	14.801	2.210	229	Perseus-Pisces (A347)
78.....	01 26 33.5	-01 34 42	29 (0)	70.73	5293	442.7	0.98	14.321	14.500	1.980	95	A194
56.....	01 09 52.6	+32 43 06	29 (0)	68.38	5064	368.8	1.30	14.286	14.442	1.848	79	
807.....	13 48 47.9	-30 18 54	28 (0)	64.59	4596	432.8	1.20	14.391	14.387	1.791	956	Centaurus (A3574)
763.....	13 17 41.1	-16 36 08	28 (0)	33.44	2332	496.4	1.09	14.469	14.900	2.800	852	
432.....	07 21 47.7	-30 04 19	28 (4)	28.69	2054	204.9	2.19	14.004	14.221	2.287	391	NGC 2280 cluster
282.....	04 22 27.8	+36 43 26	28 (0)	82.20	6032	266.6	1.79	14.145	14.255	1.560	292	
732.....	12 56 09.5	-13 37 44	27 (0)	65.56	4510	319.2	1.34	14.175	14.529	1.967	883	
295.....	04 35 49.4	-58 57 46	26 (0)	15.84	1226	188.9	1.50	13.769	14.019	2.117	391	Dorado

NOTES.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds. This catalog has been produced using parameters (D_0, V_0) = (0.89 Mpc, 350 km s⁻¹), corresponding to the density contrast $\delta\rho/\rho = 80$. We assume $h = 0.73$ where a value is required. Col. (4): Number of group members (including those generated from the population of the plane). The number derived from the Galactic-plane population is contained in parentheses. Col. (5): Mean (corrected) group distance. Col. (6): Mean heliocentric group velocity. Col. (7): Line-of-sight velocity dispersion. Col. (8): Projected virial radius. Col. (9): Log of the virial mass in solar units. Col. (10): Log of the projected mass in solar units. Col. (11): Log of the (projected) mass-to-light ratio in solar units. Col. (12): Group number from LDC catalog that encompasses all members of this group. Col. (13): Composition of group based on known galaxy clusters and superclusters.

TABLE 10
GROUPS IN THE LDC CATALOG AND THEIR MEMBERS

Name (1)	R.A. (2)	Decl. (3)	V_h (km s ⁻¹) (4)	m_K (5)	Distance (Mpc) (6)	Corresponding Group Number (7)
Group 1						
000009.14+3244182	00 00 09.0	+32 44 18	10372	10.61	140.48	1
000028.80+3246563	00 00 28.8	+32 46 56	9803	11.09	132.77	1
235950.52+3242086	23 59 50.5	+32 42 09	10086	11.12	136.61	1
Group 2						
000104.78+0432261	00 01 04.7	+04 32 26	9151	11.09	123.42	None
000221.65+0405230	00 02 21.6	+04 05 23	8606	11.10	116.03	None
235858.87+0338045	23 58 58.8	+03 38 04	8854	10.93	119.43	None
Group 3						
000037.94+2823041	00 00 37.8	+28 23 04	8705	10.46	117.75	1258
000046.96+2824071	00 00 46.8	+28 24 07	8764	10.41	118.55	1258
000433.73+2818059	00 04 33.6	+28 18 06	8785	10.62	118.79	None
235828.41+2802025	23 58 28.3	+28 02 03	9145	10.94	123.72	1258
235943.72+2817251	23 59 43.6	+28 17 25	9073	10.71	122.74	1258

TABLE 10—*Continued*

Name (1)	R.A. (2)	Decl. (3)	V_h (km s ⁻¹) (4)	m_K (5)	Distance (Mpc) (6)	Corresponding Group Number (7)
Group 4						
000100.43+0614312	00 01 00.3	+06 14 31	5324	10.03	71.77	None
000348.85+0728429	00 03 48.9	+07 28 43	5241	9.67	70.63	None
000649.47+0837425	00 06 49.5	+08 37 42	5257	10.58	70.82	None
235651.54+0530303	23 56 51.6	+05 30 30	5405	11.05	72.90	None
Group 5						
000105.97–5359303	00 01 06.0	–53 59 30	9423	10.57	129.47	None
000310.64–5444562	00 03 10.6	–54 44 56	9767	10.34	134.20	None
000311.27–5444588	00 03 11.3	–54 44 59	9767	10.35	134.20	None
Group 6						
000329.22+2721063	00 03 29.1	+27 21 06	7690	11.02	103.97	4
000548.43+2726579	00 05 48.3	+27 26 58	7531	10.95	101.81	4
000717.10+2740421	00 07 17.1	+27 40 42	7550	11.13	102.05	4

NOTES.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds. The LDC group catalog was created using $(D_0, V_0) = (1.63 \text{ Mpc}, 399 \text{ km s}^{-1})$, corresponding to a density contrast $\delta\rho/\rho = 12$. Col. (4): Heliocentric velocity. Col. (5): Corrected distance, assuming $h = 0.73$. Col. (6): Apparent K magnitude. Col. (7): Corresponding group number assigned to this galaxy when in the HDC catalog. Table 10 is published in its entirety in the electronic edition of the *Astrophysical Journal*. A portion is shown here for guidance regarding its form and content.

TABLE 11
GROUPS IN THE HDC CATALOG AND THEIR MEMBERS

Name (1)	R.A. (2)	Decl. (3)	V_h (km s ⁻¹) (4)	m_K (5)	Distance (Mpc) (6)	Corresponding Group Number (7)
Group 1						
000009.14+3244182	00 00 09.0	+32 44 18	10372	10.61	140.48	1
000028.80+3246563	00 00 28.8	+32 46 56	9803	11.09	132.77	1
235950.52+3242086	23 59 50.5	+32 42 09	10086	11.12	136.61	1
Group 2						
000001.68+4716282	00 00 01.7	+47 16 28	5017	9.68	68.91	11
000012.95+4657543	00 00 13.0	+46 57 54	5366	10.84	73.57	11
000426.65+4729250	00 04 26.6	+47 29 25	5269	10.52	72.27	11
000527.96+4632371	00 05 28.0	+46 32 37	4971	11.02	68.21	11
000723.79+4702265	00 07 23.8	+47 02 27	5313	9.93	72.81	11
000724.58+4659195	00 07 24.6	+46 59 20	5097	10.87	69.91	11
235247.40+4648138	23 52 47.3	+46 48 14	5047	10.65	69.35	11
235401.14+4729225	23 54 01.1	+47 29 22	5202	10.39	71.45	11
235526.14+4716485	23 55 26.1	+47 16 49	5348	10.98	73.38	11
235915.79+4653213	23 59 15.8	+46 53 21	5021	9.46	68.95	11
Group 3						
000126.77+3126016	00 01 26.7	+31 26 02	4948	10.23	67.25	9
000130.05+3126306	00 01 30.0	+31 26 31	4767	10.47	64.83	9
000308.87+3102108	00 03 08.9	+31 02 11	4797	11.16	65.20	9
000424.49+3128193	00 04 24.5	+31 28 19	4958	11.14	67.36	9
Group 4						
000329.22+2721063	00 03 29.1	+27 21 06	7690	11.02	103.97	6
000548.43+2726579	00 05 48.3	+27 26 58	7531	10.95	101.81	6
000717.10+2740421	00 07 17.1	+27 40 42	7550	11.13	102.05	6

TABLE 11—*Continued*

Name (1)	R.A. (2)	Decl. (3)	V_h (km s ⁻¹) (4)	m_K (5)	Distance (Mpc) (6)	Corresponding Group Number (7)
Group 5						
000457.78+0507245	00 04 57.8	+05 07 24	5357	11.19	72.17	7
000527.66+0513204	00 05 27.6	+05 13 20	5294	10.08	71.31	7
000640.35+0506483	00 06 40.3	+05 06 48	5371	11.22	72.34	7

NOTES.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds. The HDC group catalog was created using $(D_0, V_0) = (0.89 \text{ Mpc}, 350 \text{ km s}^{-1})$, corresponding to a density contrast $\delta\rho/\rho = 80$. Col. (4): Heliocentric velocity. Col. (5): Corrected distance, assuming $h = 0.73$. Col. (6): Apparent K magnitude. Col. (7): Corresponding group number assigned to this galaxy when in the LDC catalog. Table 11 is published in its entirety in the electronic edition of the *Astrophysical Journal*. A portion is shown here for guidance regarding its form and content.

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